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Analysis and calibration of Ca II triplet spectroscopy of red giant branch stars from VLT/FLAMES observations

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ABSTRACT

We demonstrate that low-resolution Ca II triplet (CaT) spectroscopic estimates of the overall metallicity ([Fe/H]) of individual red giant branch (RGB) stars in two nearby dwarf spheroidal galaxies (dSphs) agree to ± 0.1 – 0.2 dex with detailed high-resolution spectroscopic determinations for the same stars over the range $-2.5 < [\text{Fe}/\text{H}] < -0.5$. For this study, we used a sample of 129 stars observed in low- and high-resolution modes with VLT/FLAMES in the Sculptor and Fornax dSphs. We also present the data-reduction steps we used in our low-resolution analysis and show that the typical accuracy of our velocity and CaT [Fe/H] measurement is $\sim 2 \text{ km s}^{-1}$ and 0.1 dex, respectively. We conclude that CaT–[Fe/H] relations calibrated on globular clusters can be applied with confidence to RGB stars in composite stellar populations over the range $-2.5 < [\text{Fe}/\text{H}] < -0.5$.

Key words: techniques: spectroscopic – stars: abundances – galaxies: dwarf – galaxies: star clusters.

1 INTRODUCTION

An important aspect for a full understanding of galactic evolution is the metallicity distribution function of the stellar population with time.

Carrying out detailed abundance analyses with high-resolution (HR) spectroscopy to trace the patterns that allow one to distinguish between the different galactic chemical enrichment processes is time-consuming for large samples of individual stars in a galaxy. This is partly due to the observing time required, but also because of the complex data reduction and analysis necessary. Fortunately, there is an empirically developed, simply calibrated method available which can make an efficient estimate of metallicity ([Fe/H]) for individual red giant branch (RGB) stars using the strength of the Ca II triplet (CaT) lines at 8498, 8542 and 8662 Å. This method was pioneered for use on individual stars by Armandroff & Da Costa (1991). It has the advantage that the lines are broad enough that they can be accurately measured with moderate spectral resolution (e.g. Cole et al. 2004, hereafter C04).

The CaT method is routinely used to estimate [Fe/H] for nearby resolved stellar systems and also provides an accurate radial velocity estimate. Both measurements are facilitated by the strength of the

CaT lines and by the generally red colours of the target stars. However, the CaT-derived abundances are empirically defined, with a poorly understood physical basis. Therefore, it is important to check the results against HR spectroscopic (i.e. direct) measurements of [Fe/H] and other elements. The ‘classical’ CaT calibration is based on the use of globular cluster stars, all of which are drawn from a single age and metallicity stellar population. The CaT equivalent widths (EWs) are directly compared to HR spectroscopic measurements of [Fe/H] over a range of metallicity, and this comparison is used to define the relation between CaT EW and [Fe/H] for all observations taken with the same set-up. This approach has been extensively tested for a large sample of globular clusters (Rutledge, Hesser & Stetson 1997a; Rutledge et al. 1997b, and references therein). However, globular clusters typically exhibit a constant [Ca/Fe] for a large range of [Fe/H]. This leads to uncertainty in the effect of varying [Ca/Fe] ratios such as is seen in the more complex stellar populations found in galaxies. Furthermore, stars in dwarf galaxies invariably cover a significant range of ages as well as metallicities. This mismatch in the properties of calibrators and targets has led to suggestions that the CaT method may not be a very accurate indicator of [Fe/H] for more complex stellar populations, especially in those cases where [Ca/Fe] varies significantly (e.g. Pont et al. 2004).

In this paper, we investigate the validity of the CaT method for complex stellar populations. We compare large samples of [Fe/H]

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measurements coming from VLT/FLAMES made using both the CaT method and direct HR spectroscopic measurements *for the same stars* in two nearby dwarf spheroidal galaxies (dSphs), Sculptor and Fornax, over a range of [Fe/H] and [Ca/H]. This is for the first time that such a detailed comparison has been made for stars outside globular clusters. We also investigate the theoretically predicted behaviour of the CaT method for a range of stellar atmospheric parameters using a grid of model atmosphere spectra from Munari et al. (2005).

This paper is organized as follows. In Section 2, we describe the data-reduction steps we use within the Dwarf galaxy Abundances and Radial velocities Team (DART) collaboration to estimate EWs and velocities from observations in the CaT region, as the accuracy with which this can be done clearly has important implications for the reliability of our conclusions for these galaxies. We also discuss the verification of the overall calibration and accuracy of the velocity and EW measurements by comparison of results from independent CaT observations and by comparing with theoretical expectations based on the signal-to-noise ratio (S/N), resolution and line profile properties. In Section 3, we derive the standard CaT–[Fe/H] globular cluster calibration for low-resolution (LR) VLT/FLAMES data. In Section 4, we compare the derived [Fe/H] from the CaT to the HR [Fe/H] for the Sculptor and Fornax dSphs. Finally, in Section 5 we discuss the uncertainties that come from using Ca II lines to derive an [Fe/H] abundance for stellar populations where the α -abundance varies and use a comparison with stellar model atmospheres to further investigate age, metallicity and α -abundance effects.

2 LOW-RESOLUTION DATA REDUCTION AND ANALYSIS

The data sets presented here were collected between 2003 August and 2005 November. They consist of 15 pointings in the Sculptor dSph and 11 in the Fornax dSph spread over the galaxies (Fig. 1). Some fields were observed with 1 h exposure time, whilst other fields have repeated exposures of shorter integration time and two different plate set-ups, with the aim of testing the reliability of the derived velocities, EWs and the stability of the instrument.

All the LR CaT observations were made using VLT/FLAMES in Medusa mode. This allows the simultaneous allocation of up to 132 fibres, including dedicated sky fibres, over a 25 arcmin diameter field of view. We used the GIRAFFE LR grating (LR8), which covers the wavelength range from 8206 to 9400 Å, and gives a resolution of $R \approx 6500$. This allows the measurement of EWs from CaT lines and also enables the derivation of velocities accurate to approximately a few km s^{−1}. This set-up was used as part of the DART programme to obtain spectra for several different fields in each of the Sculptor, Fornax and Sextans dSphs, and also for calibration purposes on a sample of four globular clusters: NGC 104, NGC 5904, NGC 3201 and NGC 4590,¹ which cover the range $-2.0 \lesssim [\text{Fe}/\text{H}] \lesssim -0.7$ on the Carretta & Gratton (1997) (hereafter CG97) scale (see Table 1).

The central fields in Sculptor and Fornax were also observed with a similar VLT/FLAMES set-up but at HR with $R \approx 20\,000$, which facilitates direct measurement of individual lines, and hence direct abundance determination, of numerous elements (see Section 4).

Table 2 shows the journal of the VLT/FLAMES LR and HR observations we used for our analysis of the CaT–[Fe/H] calibration (Sections 3 and 4).

¹ We acquired the data for NGC 104 during an observing run in 2005 January, whilst the data for the other globular clusters are from the ESO archive.

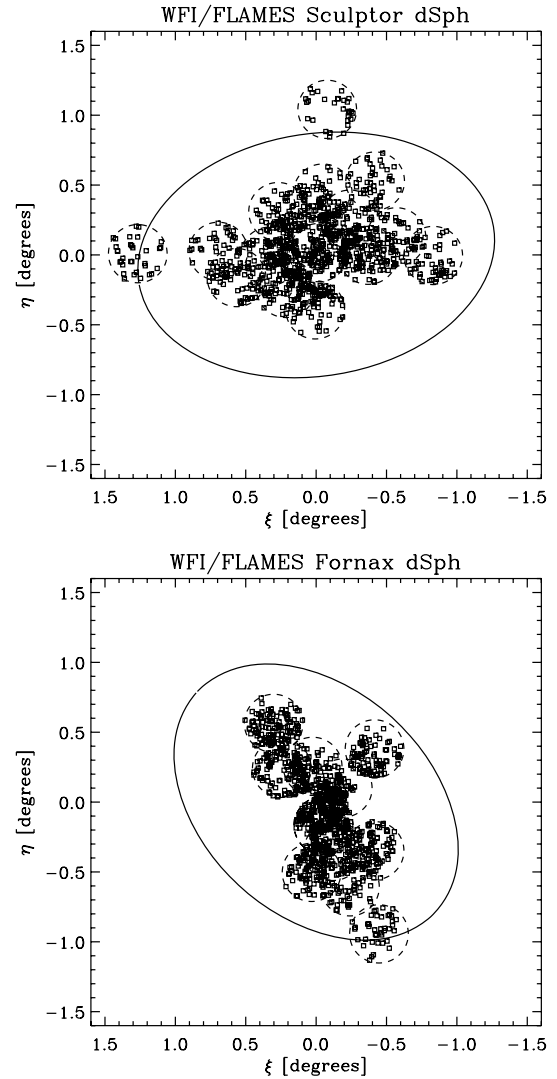


Figure 1. Locations of the observed FLAMES targets at LR (squares) for the Sculptor dSph (top panel), centred at $01^{\text{h}}00^{\text{m}}09^{\text{s}}$, $-33^{\circ}42'30''$ (Mateo 1998) with a position angle 99° (Irwin & Hatzidimitriou 1995), and for the Fornax dSph (bottom panel), centred at $02^{\text{h}}39^{\text{m}}52^{\text{s}}$, $-34^{\circ}30'49''$ with a position angle $46^{\circ}8$ (Battaglia et al. 2006). The ellipses indicate the tidal radius (Sculptor: Irwin & Hatzidimitriou 1995; Fornax: Battaglia et al. 2006). The dashed circles show the FLAMES pointings. North is up and east is to the left-hand side.

The data were all initially reduced using the GIRBLDRS² pipeline provided by the FLAMES consortium (Geneva Observatory, Blecha et al. 2003). This package provided flat-fielding (including fringing removal), individual spectral extraction and accurate wavelength calibration, based on daytime calibration exposures. At the time we started this project, no sky subtraction was available within this pipeline which led us to develop several further reduction stages for the LR analysis. We describe the LR analysis in detail here. In Tables 3 and 4, we present the LR and HR results relevant for our analysis of the CaT–[Fe/H] calibration (Sections 3–5).

² Available at SourceForge, <http://girbldrs.sourceforge.net/>.

Table 1. The globular clusters used in this work. Positions and V_{HB} are from Harris (1996). The systemic velocity v_{sys} was derived in this work. $[Fe/H]_{CG97}$ is the value of metallicity from CG97. $[Fe/H]_{R97}$, $[Fe/H]_{T01}$ and $[Fe/H]_{C04}$ are the values of the metallicity from this work, using the calibrations from Rutledge et al. (1997a,b), T01 and C04, that is, equations (13), (14) and (15), respectively. The errors in the derived metallicities were obtained from the standard deviation of the values of reduced EW for each globular cluster. The $[Ca/Fe]$ values are from Carretta et al. (2004) for 47 Tuc and all the other ones are listed in Koch et al. (2006).

Cluster	RA(J2000)	Dec.(J2000)	V_{HB}	v_{sys} (km s $^{-1}$)	$[Fe/H]_{CG97}$	$[Fe/H]_{R97}$	$[Fe/H]_{T01}$	$[Fe/H]_{C04}$	$[Ca/Fe]$
NGC 104	00 ^h 24 ^m 05 ^s .2	−72°04′51″	14.06	−14.4 ± 0.7	−0.70 ± 0.03	−0.78 ± 0.05	−0.68 ± 0.06	−0.82 ± 0.06	0.30
NGC 5904	15 ^h 18 ^m 33 ^s .8	+02°04′58″	15.07	53.2 ± 0.7	−1.11 ± 0.03	−1.13 ± 0.07	−1.02 ± 0.08	−1.19 ± 0.06	0.21
NGC 3201	10 ^h 17 ^m 36 ^s .8	−46°24′40″	14.77	497.5 ± 1.2	−1.23 ± 0.05	−1.27 ± 0.07	−1.17 ± 0.06	−1.36 ± 0.07	0.11
NGC 4590	12 ^h 39 ^m 28 ^s .0	−26°44′34″	15.68	−98.3 ± 1.7	−1.99 ± 0.06	−1.94 ± 0.06	−1.90 ± 0.07	−2.11 ± 0.06	0.32

Table 2. Table of VLT/FLAMES observations for the Sculptor and Fornax dSphs and the globular clusters NGC 104, NGC 3201, NGC 4590 and NGC 5904. We used these observations for the analysis of the CaT– $[Fe/H]$ calibration.

Object	RA(J2000)	Dec.(J2000)	Date of observation (yy-mm-dd)	Exposure time (s)	Grating
Sculptor	01 00 03	−33 41 30	2003-08-20	4200	HR13
Sculptor	01 00 03	−33 41 30	2003-08-21	3 × 3600, 4700	HR14
Sculptor	01 00 03	−33 41 30	2003-08-22	2 × 3600	HR13
Sculptor	01 00 03	−33 41 30	2003-08-22	4500	HR14
Sculptor	01 00 03	−33 41 30	2003-08-23	5400	HR14
Sculptor	01 00 03	−33 41 30	2003-08-24	2 × 3600	HR10
Sculptor	01 00 03	−33 41 30	2003-08-25	5400	HR10
Sculptor	01 00 03	−33 41 30	2003-08-28	3600	HR10
Sculptor	00 58 22	−33 42 36	2003-09-29	2 × 1800	LR8
Sculptor	01 01 44	−33 43 01	2003-10-01	2 × 1800	LR8
Sculptor	01 00 03	−33 41 29	2004-09-10	3600	LR8
Sculptor	01 01 11	−33 56 50	2005-11-07	3600	LR8
Sculptor	00 58 59	−33 27 04	2005-11-08	2 × 2700	LR8
Fornax	02 39 30	−34 28 39	2003-09-29	2 × 3600	HR13
Fornax	02 39 30	−34 28 39	2003-09-29	2 × 3100	HR14
Fornax	02 39 30	−34 28 39	2003-09-29	2 × 3600	HR13
Fornax	02 39 30	−34 28 39	2003-09-30	3300, 2 × 5400	HR14
Fornax	02 39 30	−34 28 39	2003-09-30	3600	HR10
Fornax	02 39 30	−34 28 39	2003-10-01	3 × 3600	HR10
Fornax	02 39 30	−34 28 39	2003-10-01	3300, 3600	HR14
Fornax	02 39 30	−34 28 39	2004-01-14	3600	HR10
Fornax	02 39 30	−34 28 39	2004-01-15	3600	HR13
Fornax	02 39 30	−34 28 39	2004-01-19	2 × 3600	HR13
Fornax	02 39 30	−34 28 39	2004-01-20	3600	HR13
Fornax	02 39 30	−34 28 39	2004-01-21	2 × 3600	HR14
Fornax	02 39 30	−34 28 39	2004-01-22	3600, 3900	HR14
Fornax	02 39 30	−34 28 39	2004-01-23	3600	HR10
Fornax	02 39 30	−34 28 39	2004-01-23	3600	HR14
Fornax	02 39 30	−34 28 39	2004-01-24	3600	HR10
Fornax	02 39 30	−34 28 39	2004-01-26	3600	HR10
Fornax	02 38 56	−34 24 48	2005-11-07	3600	LR8
Fornax	02 39 36	−34 40 51	2003-08-24	1500, 1800	LR8
Fornax	02 39 58	−34 15 28	2003-08-26	2400	LR8
NGC 104	00 22 15	−72 04 45	2005-01-11	3 × 600	LR8
NGC 3201	10 17 36	−46 24 40	2003-02-22	300, 327	LR8
NGC 3201	10 17 36	−46 24 40	2003-03-06	2 × 600	LR8
NGC 4590	12 39 28	−26 44 33	2003-03-06	3 × 600	LR8
NGC 5904	15 18 03	02 06 40	2003-03-05	6 × 600	LR8

2.1 Sky subtraction and wavelength calibration

An example of a spectrum produced by the GIRBLDRS pipeline is shown in Fig. 2. The numerous skylines visible in this part of the spectrum not only serve as an independent check on the overall wavelength calibration, but also enable updation of the wavelength calibration of the individual spectra.

Since the CaT lines only occupy a limited part of the LR spectra, we optimized the sky subtraction and wavelength refinement for the range 8400–8750 Å.

The first step is to combine all the sky spectra (typically 10–20 sky fibres were allocated per field) using k-sigma clipping to remove spurious features and obtain an average sky spectrum. The result is then split into continuum and sky-line components using

Table 3. This table lists a sample of the relevant data for the stars in the globular clusters NGC 104, NGC 3201, NGC 4590 and NGC 5904 observed with VLT/FLAMES at LR, and that we used in the analysis of the CaT–[Fe/H] calibration. The full table is available as Supplementary Material to the online version of this article. For the CaT analysis, we select only those stars whose colour and magnitude are consistent with what is expected for RGB stars, and that have $V - V_{\text{HB}} < 0$, $S/N > 10$ per Å, error in velocity $< 5 \text{ km s}^{-1}$ and velocities consistent with membership to the cluster (we assign membership to those stars within 3σ of the systemic velocity, see Table 1). The columns indicate: (1): the globular cluster name; (2): the star ID; (3) and (4): star coordinates (right ascension in hours and declination in degrees); (5) and (6): V magnitude and its error; (7) and (8): heliocentric velocity and its error; (9) and (10): summed CaT EW ($\text{EW}_2 + \text{EW}_3$) and its error; and (11) and (12): [Fe/H] value and its error, derived from the LR observations applying equation (16). The photometry and astrometry are from Stetson (2000).

Object	Star ID	RA(J2000)	Dec.(J2000)	V	eV	v_{hel}	ev_{hel}	ΣW_{T01}	$e\Sigma W_{\text{T01}}$	[Fe/H] _{LR}	$e[\text{Fe}/\text{H}]_{\text{LR}}$
NGC 104	GB07-TUC1002	0 23 41.71	−72 5 2.6	13.052	0.002	−8.55	0.38	5.209	0.055	−0.80	0.02
NGC 104	GB07-TUC1510	0 24 25.01	−72 2 45.7	12.652	0.001	−6.93	0.57	5.729	0.039	−0.69	0.02
NGC 3201	GB07-N3201-S60	10 17 42.84	−46 22 48.5	13.601	0.002	495.55	0.65	4.144	0.055	−1.32	0.02
NGC 3201	GB07-N3201-S74	10 17 48.53	−46 23 39.3	14.602	0.002	498.98	0.63	3.636	0.061	−1.26	0.03
NGC 4590	GB07-M68-S99	12 39 40.21	−26 47 37.5	15.186	0.004	−100.62	1.80	2.021	0.132	−2.06	0.06
NGC 4590	GB07-M68-S144	12 39 32.06	−26 47 53.6	15.291	0.004	−99.36	1.92	2.139	0.139	−1.98	0.06
NGC 5904	GB07-M5-S23	15 17 55.26	2 7 25.6	14.544	0.002	55.11	1.08	4.175	0.036	−1.12	0.02
NGC 5904	GB07-M5-S980	15 18 16.42	2 7 18.5	14.264	0.002	46.88	1.16	4.525	0.023	−1.05	0.01

an iterative k -sigma clipped non-linear filter (a combination of a median and a boxcar). The average ‘sky-line’ spectrum is then used to define a sky-line template mask, in order to isolate those regions of the sky spectrum with significant features and mask out the remainder.

The processing of individual object spectra then proceeds as follows.

(i) Each spectrum is filtered as above to split the spectra into a line and continuum component, but this time additionally masking out those regions affected by sky lines to allow a more accurate definition of the continuum.

(ii) The object line spectrum, which includes sky lines, is then cross-correlated with the masked line component of the average sky spectrum. This provides an accurate differential wavelength update. The object line spectrum is then (re)interpolated to be on the same wavelength scale as the average sky spectrum.

(iii) For the sky subtraction, we compare the masked sky-line and object-line spectra, and find the optimum scalefactor and profile-matching kernel that produce the minimum average absolute deviation (L1 norm). This is applied to the line-only spectra.

The optimum scaling factor, which in this case is chosen to minimize the L1 norm rather than the commoner L2 norm to reduce sensitivity to non-Gaussian outliers, is derived using a simple grid search with progressively finer step size. As noted previously, a mask is used to isolate the relevant regions of the sky spectrum to match to a template. By first removing the continua from both sky and object, more emphasis is placed on minimizing the impact of sky-line residuals.

(iv) Finally, the object continuum is added back to the wavelength-updated object-line spectrum, the sky continuum is removed using the sky-line scaling factor and the sky-subtracted spectrum is saved for the next stage in the processing. Of course, implicit in the sky correction is the reasonable assumption that the derived scalefactor for both sky lines and sky continuum is the same.

The sky-subtraction process described above involves some key components which are crucial to achieve good results. Accurate wavelength registration is absolutely vital for good sky subtraction and is facilitated by the presence of copious numbers of strong sky lines. These sky lines are unresolved at this resolution and at the S/N achieved (see later) which readily enables sub- km s^{-1} precision in wavelength alignment. This also has the added advantage of

ensuring that systematic offsets due to wavelength calibration for velocities from different observations are negligible.

As a final step in this process, all the average sky spectra from each FLAMES observation, are cross-correlated with a chosen reference sky spectrum and used to put all the observations on the same internal system. This is done to avoid possible systematic differences between observations taken at different times.

The effects of combining the sky spectra to form an average sky and re-interpolating the object spectrum to this average wavelength system, almost invariably results in a slight mismatch between the spectral line profiles of the object and average sky spectra. This is circumvented by applying a Hanning smoothing kernel to each in turn and finding which combination of smoothed and unsmoothed gives the best results (as determined by the optimum scalefactor). More sophisticated adaptive kernel matching (e.g. Alard & Lupton 1998) is probably unwarranted in this case.

2.2 Velocity and EW estimation

Our goal is to produce a robust automatic procedure that gives close to optimum results in terms of the S/N and also produces minimal systematic bias as a function of EW.

The first stage of the process is to estimate the continuum in a similar manner to that described in the previous section. Each spectrum is split into a line-only, and smoothly varying continuum-only component, using an iterative k -sigma clipped non-linear filter. This time the region around each CaT line is masked out, prior to filtering, to prevent the continuum tracking the wings of strong CaT lines. The effective scalelength of the filtering is set to $\sim 15 \text{ Å}$ which, in conjunction with the masking and iterative clipping, is sufficient to follow continuum trends without being affected by the presence of strong lines. At the same time, an estimate of the overall S/N in the continuum is made by measuring the median continuum level in the CaT region and the pixel-to-pixel noise covariance matrix in regions containing no lines. The latter is needed to correct the apparent random noise for the cumulative smoothing effects of spectral interpolation and resampling, which we find typically results in a factor of ≈ 2 overall random noise reduction.

After normalizing by the computed continuum, the velocity is estimated by cross-correlating each spectrum with a template. This template is constructed from a zero-level continuum superimposed on three Gaussian absorption lines located at the vacuum rest

Table 4. This table lists a sample of the relevant data for the stars observed with VLT/FLAMES at both LR and HR in the Sculptor and Fornax dSphs, and used in the analysis of the CaT–[Fe/H] calibration. The full table is available as Supplementary Material to the online version of this article. The columns indicate: (1): the galaxy name; (2): the star ID; (3) and (4): star coordinates (right ascension in hours and declination in degrees); (5) and (6): V magnitude and its error; (7) and (8): heliocentric velocity and its error; (9) and (10): summed CaT EW ($EW_2 + EW_3$) and its error; (11) and (12): [Fe/H] value derived from the LR observations applying equation (16) and its error; (13) and (14): HR [Fe/H] value and error; (15) and (16): [Ca/Fe] and its error, from the HR observations. The photometry and astrometry are from our ESO/WFI observations (Tolstoy et al. 2004; Battaglia et al. 2006). Those stars with [Ca/Fe] = -9.99 had too low S/N to allow a determination of the HR [Ca/H].

Object	Star ID	RA(J2000)	Dec(J2000)	V	eV	v_{hel}	$e v_{\text{hel}}$	ΣW_{TOT}	$e \Sigma W_{\text{TOT}}$	[Fe/H] _{LR}	$e[\text{Fe}/\text{H}]_{\text{LR}}$	[Fe/H] _{HR}	$e[\text{Fe}/\text{H}]_{\text{HR}}$	[Ca/Fe] _{HR}	$e[\text{Ca}/\text{Fe}]_{\text{HR}}$
Sculptor	GB07-Sc01	1 0 2.69	-33 30 25.3	18.017	0.003	109.24	1.31	4.269	0.250	-1.53	0.11	-1.39	0.10	0.10	0.12
Sculptor	GB07-Sc02	1 0 25.78	-33 30 25.4	17.692	0.002	119.62	1.49	3.392	0.267	-2.00	0.12	-1.93	0.11	0.30	0.13
Sculptor	GB07-Sc03	0 59 37.00	-33 30 28.4	17.706	0.002	101.70	1.01	4.344	0.087	-1.58	0.04	-1.63	0.10	0.27	0.13
Sculptor	GB07-Sc04	0 59 55.63	-33 33 24.6	16.890	0.002	120.51	1.13	3.793	0.121	-2.05	0.05	-1.88	0.10	0.23	0.13
Sculptor	GB07-Sc05	0 59 47.21	-33 33 36.9	17.608	0.002	105.13	0.85	3.362	0.162	-2.04	0.07	-2.25	0.10	0.37	0.13
Sculptor	GB07-Sc06	1 0 21.08	-33 33 46.4	18.082	0.003	107.34	1.94	4.755	0.221	-1.30	0.10	-1.30	0.10	0.17	0.13
Fornax	GB07-Fnx01	2 39 22.15	-34 19 40.4	18.355	0.004	68.45	0.88	4.155	0.201	-1.81	0.09	-1.44	0.05	0.02	0.10
Fornax	GB07-Fnx02	2 39 51.42	-34 21 21.0	18.548	0.005	59.97	2.08	5.869	0.217	-1.00	0.10	-0.79	0.05	-0.44	0.07
Fornax	GB07-Fnx03	2 39 31.49	-34 23 5.1	18.534	0.005	53.72	0.88	5.564	0.150	-1.14	0.07	-0.85	0.05	-0.17	0.05
Fornax	GB07-Fnx04	2 38 49.28	-34 24 5.0	18.543	0.004	55.20	1.61	6.097	0.145	-0.90	0.06	-0.93	0.06	-0.22	0.07
Fornax	GB07-Fnx05	2 39 4.31	-34 25 18.9	18.390	0.004	55.64	1.89	5.275	0.154	-1.31	0.07	-0.95	0.04	-0.11	0.06
Fornax	GB07-Fnx06	2 38 55.53	-34 25 36.3	18.188	0.004	91.68	1.07	2.945	0.135	-2.39	0.06	-2.65	0.04	0.36	0.06

wavelengths of the CaT lines. Most of the weight in the least-squares fit comes from the core of the line which is sufficiently Gaussian-like in LR data, that a Gaussian fit provides an estimate with effectively a minimal rms error. Using more complex line profiles with correspondingly greater numbers of free parameters generally makes the rms error worse and is also more prone to wildly unstable solutions due to the inevitable presence of occasional artefacts in the data. The Gaussian line depths are scaled in the ratio 3:5:4 to reflect the true relative strengths of CaT lines and all are set to have a full width at half-maximum (FWHM) = 2.35 Å. An example of a continuum fitting to a Sculptor dSph K-giant spectrum together with the computed cross-correlation function and Gaussian fit to the peak is shown in Fig. 3.

An accurate estimate of the position of the cross-correlation peak is made by fitting a Gaussian to a localized region around the peak. This velocity is then used to define the wavelength region around each CaT line to use for EW estimation. We estimate the EW in two ways. The first consists in simply summing the flux contained in a region centred on each CaT line. After some trial and error, we settled on a region 15 Å wide centred on each line as a reasonable trade-off between including all the line flux and minimizing the noise. To derive the second estimate, we fit individual unconstrained Gaussian functions to each CaT line over the same wavelength region (see Fig. 3). This also allows a semi-independent check on the accuracy of the derived velocity by providing three separate velocity measures with associated errors. The weighted sum of these velocity errors provides the basic error estimate for the velocity derived from the cross-correlation. We prefer to use the latter method for the velocity estimate since it is effectively a constrained model fit. As a final step, the derived velocity is corrected to the heliocentric system.

The combined EWs for CaT lines #2 and #3 (λ_{8542} , λ_{8662}) for both the integral and Gaussian fits are then compared and used to compute an overall correction to the Gaussian fit. This is necessary since the real line profile is a complex function of many parameters, and in particular, the dampening wings visible in strong lines are distinctly non-Gaussian in appearance. This means that the observed CaT lines have non-Gaussian wings which are progressively more visible as the EW increases. To compensate for this, we compare the ensemble Gaussian fit to the conventional EW integration for each data set analysed, by finding the best-fitting slope that links the two sets. This correction is equivalent to computing the average overlap integral between the real line profiles and the Gaussian fits and is accomplished by measuring the gradient between the two as a function of EW. By using all the data to do this, we introduce no additional rms error and remove the majority of the systematic bias as a function of EW. We also automatically take care of any generic spectrograph-induced line profile abnormalities. For the range of EWs considered here, this method works reliably (see Fig. 4). In deriving this correction, we neglect line #1 (λ_{8498}) as it is the weakest and would only add noise to the determination.

In the rest of our analysis, we use the Gaussian-derived EW estimator since although the simple integral is unbiased it is also significantly noisier than a Gaussian fit. We show in the next section how the measurements of the continuum level and the random noise error in the continuum can be used to estimate errors in the EWs and velocities.

2.3 Error bounds for velocities and EWs

For detailed abundance work most lines of interest are weak, for example, $EW \approx 100 \text{ mÅ}$. Notable exceptions are the CaT lines which are generally heavily saturated and on the damped part of

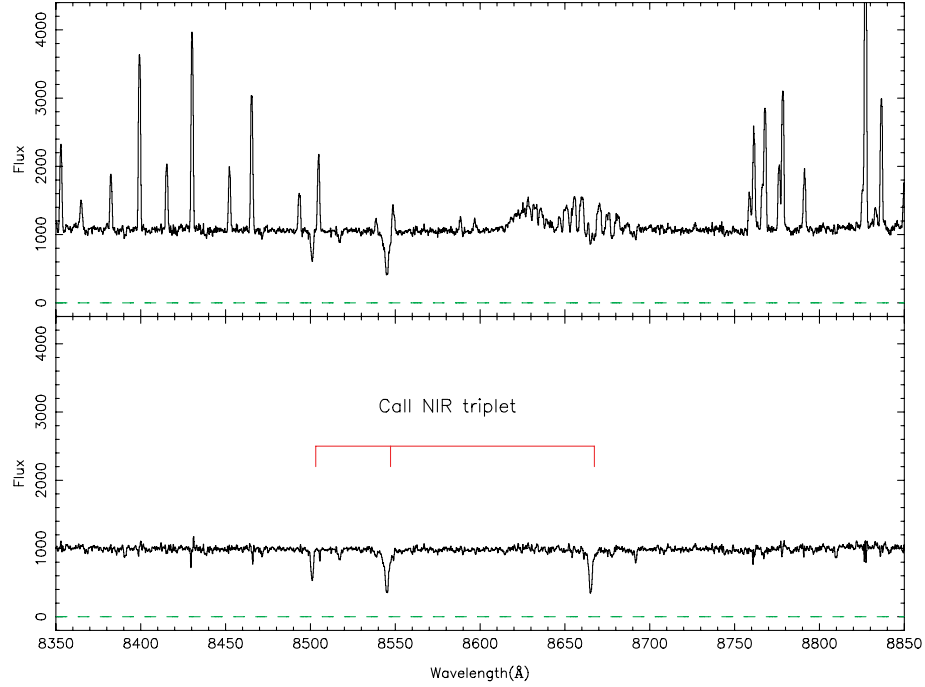


Figure 2. An example of an LR spectrum from FLAMES in the CaT wavelength range. In the upper panel is the un-sky-subtracted spectrum and in the lower panel is the result of automated sky subtraction. The CaT lines are marked.

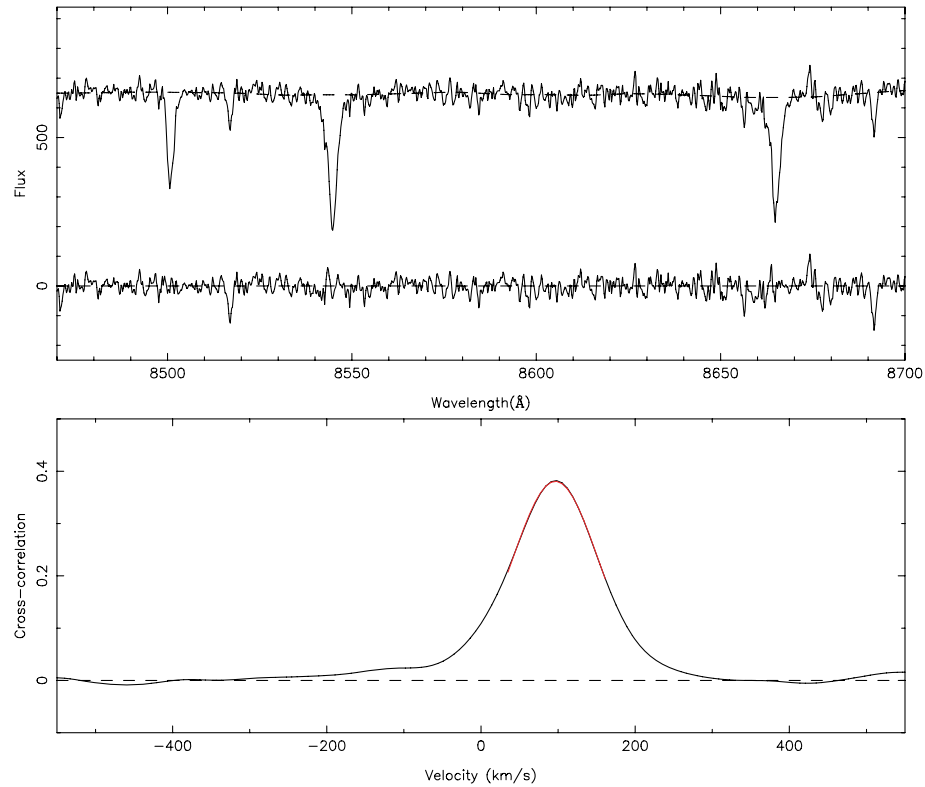


Figure 3. Top panel: the upper spectrum shows the CaT region for a K-giant in the Sculptor dSph ($V = 17.7$, $S/N \approx 40/\text{\AA}$, derived $[\text{Fe}/\text{H}]_{\text{CaT}} = -1.50$ dex), showing an example of automated continuum fitting (dashed line); the lower spectrum shows the residuals after continuum removal and subtracting Gaussian model fits to the three CaT lines. Bottom panel: the derived cross-correlation function and associated Gaussian fit around the peak region ($\pm 50 \text{ km s}^{-1}$ about the peak). The fit is so good that it is indistinguishable from the observations.

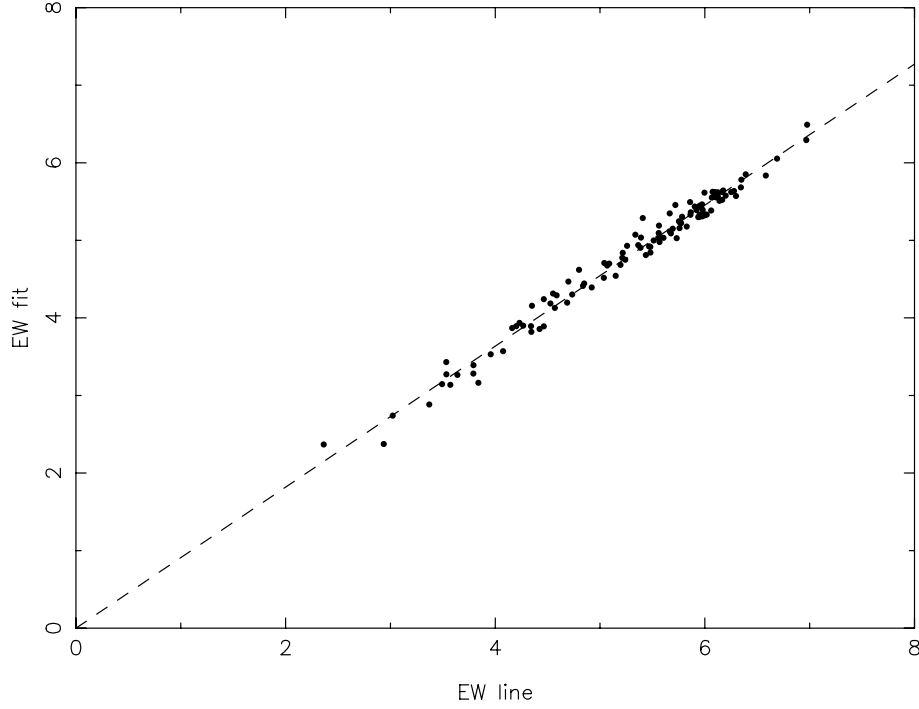


Figure 4. A comparison of the EW obtained from profile fitting (EW fit) versus the EW from the direct integral of the lines (EW line), as an example of the relationship between the two.

the curve of growth. Turbulence plus rotation of late-type giants typically only broaden the line profiles by approximately a few km s^{-1} ; hence, the profile of the CaT lines, which are completely unresolved at $R = 6500$, are dominated by intrinsic broadening due to saturation (typically $\text{FWHM} = 2\text{--}3 \text{ \AA}$) and to a lesser extent by the resolution of the spectrograph ($\text{FWHM} \approx 1.3 \text{ \AA}$).

Despite the dampening wings, to first order the LR CaT lines can be reasonably well approximated by Gaussian functions (e.g. see Fig. 3) and we can use this to gain some insight into the limiting factors that determine the accuracy of the velocity and EW measurements.

Here, we define the resolution as $\text{FWHM} = 2.35\sigma$, where σ is the Gaussian profile equivalent scale parameter. Since the total line flux is then $I_p \sqrt{2\pi}\sigma = I_p (\text{FWHM}) 1.07$, where I_p is the ‘peak’ flux, this implies that line saturation (i.e. $I_p = C$, where C is the continuum level per \AA) occurs when $\text{EW} \approx \text{FWHM}$. The intrinsic FWHM of weak lines in these late-type giants is typically only a few km s^{-1} , that is, lines with EWs of $\approx 100 \text{ m\AA}$ and above are saturated. As noted previously, the CaT lines are heavily saturated and typically have EWs well above 1 \AA .

To a reasonable approximation, the noise in the continuum (σ_n per \AA) due to sky plus object dominates, and over any individual line C can be taken to be constant over the lines of interest. Therefore, the EW and its error due to random noise, ΔEW , are given by

$$\text{EW} = \frac{\eta}{C} \quad \Delta\text{EW} = \frac{\sigma_n \sqrt{w}}{C} = \frac{\sqrt{w}}{S/N}, \quad (1)$$

where the total line flux is η , S/N is the S/N per \AA and w is the effective width (\AA) the line is integrated over.

For a fitted Gaussian profile,

$$w = \sqrt{4\pi}\sigma \approx 1.5 (\text{FWHM}), \quad (2)$$

hence

$$\Delta\text{EW} = \frac{\sqrt{1.5 (\text{FWHM})}}{S/N}. \quad (3)$$

The FWHM of the line and continuum S/N are the primary abundance error drivers from a random noise point-of-view. For example, the two strongest CaT lines, λ_{8542} and λ_{8662} , used in our analysis have FWHM at a resolution of $R = 6500$ of $2\text{--}3 \text{ \AA}$, which for a continuum S/N of 10 per \AA implies a lower bound on the combined EW error of $\approx 0.3 \text{ \AA}$.

In a similar way, we can place constraints on the accuracy of measuring velocities. For Gaussian-like line profiles, which are good approximations even for saturated lines like the CaT, the minimum variance bound on the error in the estimated line position $\hat{\lambda}$ is given by

$$\text{var}\{\hat{\lambda}\} = \frac{\sigma^2}{\eta} \frac{\sigma_n^2 \sqrt{16\pi} \sigma}{\eta}, \quad (4)$$

where the line flux is η and where the noise in the continuum, σ_n (per \AA), dominates and can be taken to be a constant over the region of interest (see Irwin 1985, for more details). Rewriting this in terms of the FWHM of the line, the EW and the continuum S/N , leads to

$$\Delta\hat{\lambda} \approx \frac{\text{FWHM}^{3/2}}{\text{EW} (S/N)}, \quad (5)$$

where all measurements are in per \AA .

As an example, for observations at $R = 6500$, the FWHM of the strongest CaT line is typically between 2 and 3 \AA , while the EW of these lines is typically $\approx 2 \text{ \AA}$, implying accuracies of wavelength centring at a continuum S/N of 10 per \AA of around 0.2 \AA per line, or equivalently $\approx 5 \text{ km s}^{-1}$ using all three CaT lines. This is the minimum S/N we consider acceptable in our analysis and is achievable on VLT FLAMES for $V = 20$ objects in 3600 s of integration.

2.4 Errors for velocity and EW from repeated measurements

The number of independent measurements for each data set is 1740 for Sculptor and 1359 for Fornax. For Sculptor, we have 464 stars observed only once, 428 stars observed twice, 73 stars observed three times, 39 observed four times and nine observed five times. For Fornax, we have 816 stars observed only once, 209 stars with double, 23 stars with triple, 14 stars with quadruple measurements.

We test for the reliability of our velocity and related errors by analysing the distribution of velocity differences from double mea-

surements. The j th observed velocity v_{ji} for a star i can be considered a random variable which follows a Gaussian distribution centred around the true value $v_{\text{true},i}$ with a dispersion given by the velocity error σ_{ji} . The difference between two repeated measurements $v_{1,i}$ and $v_{2,i}$, $\Delta v_i = v_{1,i} - v_{2,i}$, is a random variable following a Gaussian distribution centred around zero and with dispersion given by $\sigma_i = \sqrt{\sigma_{1,i}^2 + \sigma_{2,i}^2}$. Thus, if both velocities and their errors are correctly determined, the distribution of velocity differences Δv_i normalized by σ_i should be a Gaussian with mean zero and dispersion unity. Fig. 5 shows that if we take into consideration all the

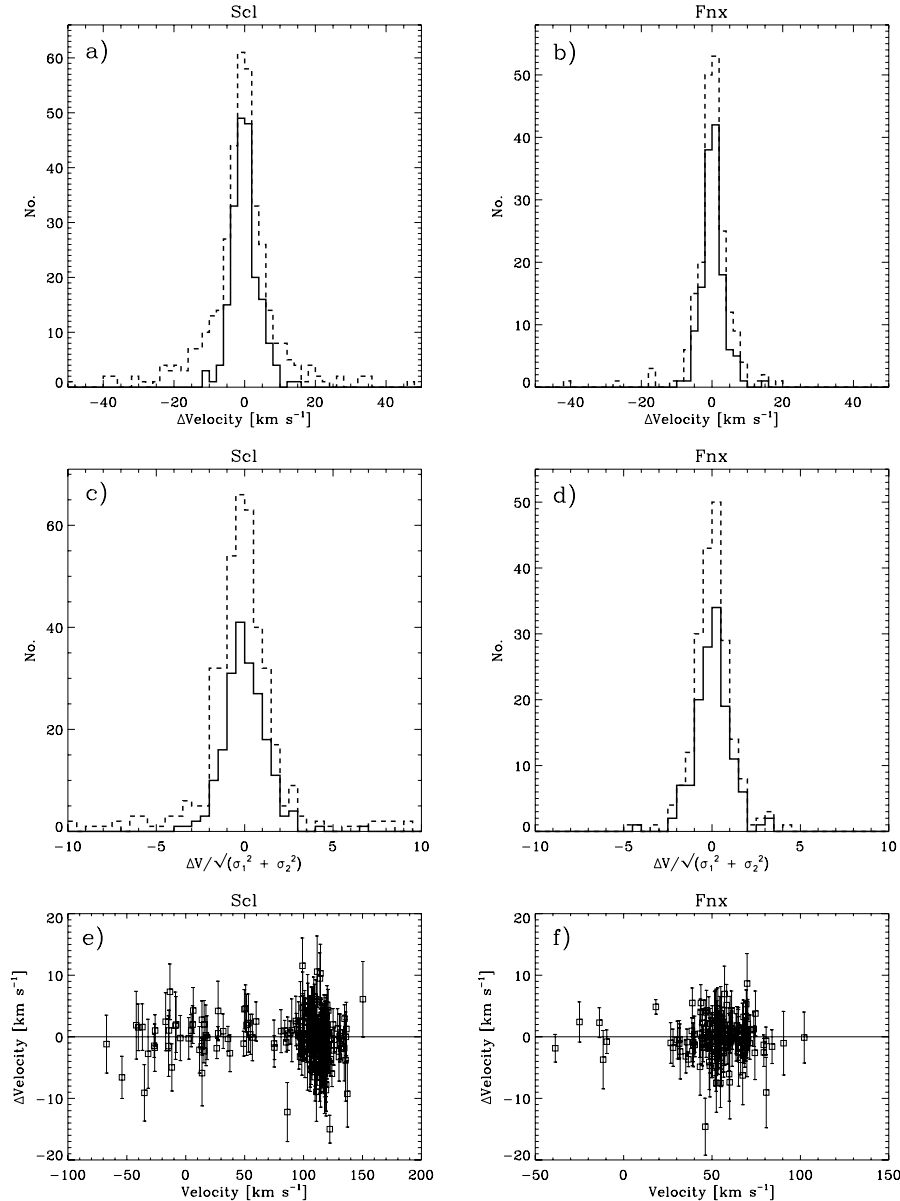


Figure 5. Comparison between velocity measurements for stars with double measurements in the Sculptor (left-hand panel) and Fornax (right-hand panel) dSphs. Panels (a) and (b): distribution of velocity differences for all the stars (dashed line, 428 stars for Sculptor and 209 for Fornax), and for the stars with S/N per Å > 10 and estimated error in velocity < 5 km s⁻¹ for each measurement (solid line, 203 stars for Sculptor and 138 for Fornax). The weighted mean velocity, rms dispersion and scaled MAD from the median (1.48 MAD ≡ a robust rms e.g. Hoaglin, Mosteller & Tukey 1983) are: -0.5 ± 1.3 km s⁻¹, 25.1 ± 1.0 km s⁻¹, 6.7 ± 0.2 km s⁻¹ (dashed line) and 0.1 ± 0.3 km s⁻¹, 3.1 ± 0.4 km s⁻¹, 3.1 ± 0.4 km s⁻¹ (solid line) for Sculptor; 1.4 ± 2.0 km s⁻¹, 27.7 ± 1.4 km s⁻¹, 3.2 ± 0.2 km s⁻¹ (dashed line) and -0.01 ± 0.36 km s⁻¹, 2.6 ± 0.4 km s⁻¹, 2.4 ± 0.4 km s⁻¹ (solid line) for Fornax. Panels (c) and (d): the same as above but now the velocity difference is normalized by the predicted error. With these S/N and velocity error cuts, the measured error in the velocity distribution is very close to the expected unit variance Gaussian (standard deviation = 1.2 and MAD = 1.1 for Sculptor; standard deviation = 1.0 and MAD = 0.9 for Fornax). Panels (e) and (f): comparison of velocities for stars with an S/N per Å > 10 and error in velocity < 5 km s⁻¹.

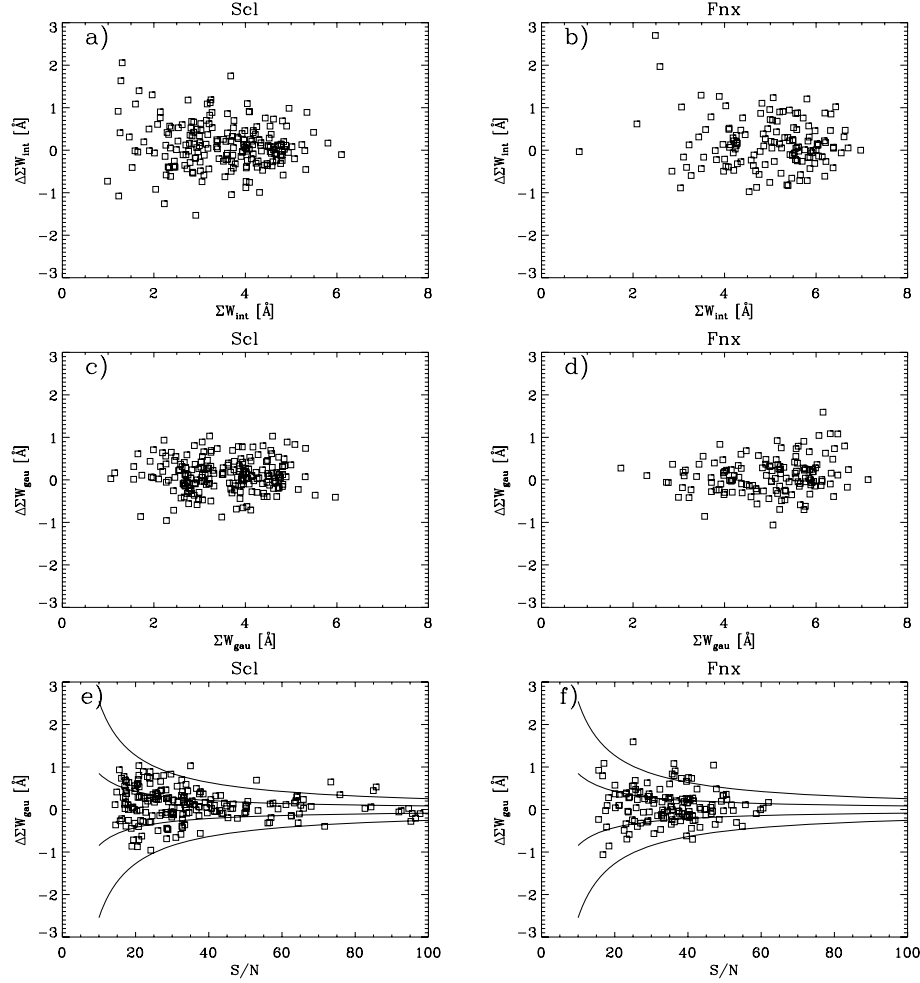


Figure 6. Comparison between summed EW ($EW_2 + EW_3$) measurements of CaT lines from integrated flux (panels a and b) and Gaussian fitting (panels c–f) for stars with double measurements, $S/N > 10$ and error in velocity $< 5 \text{ km s}^{-1}$ in the Sculptor (left-hand panel, 203 stars) and Fornax (right-hand panel, 138 stars) dSphs. The solid lines in panels (e) and (f) indicate the 1 and 3σ regions for an error in summed EW given by $\sigma_{\Sigma W} = 6/(S/N)$ (see the text). Assuming this error, the weighted average, rms and MAD for the differences in ΣW from Gaussian-derived estimator (panels c and d) is $0.04 \pm 0.04 \text{ Å}$, $0.30 \pm 0.05 \text{ Å}$, $0.32 \pm 0.05 \text{ Å}$ for Sculptor and $0.07 \pm 0.05 \text{ Å}$, $0.36 \pm 0.05 \text{ Å}$, $0.34 \pm 0.05 \text{ Å}$ for Fornax.

stars with repeated measurements both for Sculptor and Fornax the distribution has large tails of stars with $\Delta v/\sigma > 3$. If we restrict ourselves to the stars with $S/N \text{ per Å} > 10$ and error in velocity $< 5 \text{ km s}^{-1}$, then the resulting distribution is very close to a Gaussian. For Sculptor, the resulting standard deviation is 1.2 and the scaled median absolute deviation (MAD) is 1.1, with just five stars with $\Delta v/\sigma > 3$ (2 per cent). For Fornax, the standard deviation is 1.0 and the scaled MAD is 0.9.

Hereafter, we consider an $S/N \text{ per Å} > 10$ and an error in velocity $< 5 \text{ km s}^{-1}$ as the minimum requirements for a reliable determination of velocity and EW.

Figs 6(a)–(d) show the comparison of summed EW for the two strongest CaT lines, $\Sigma W = EW_2 + EW_3$, which we use as an indicator of metallicity (see Section 3), for the stars with double measurements in Sculptor and Fornax, respectively. As mentioned in Section 2.2 the Gaussian-derived estimator is considerably less noisy than the integral estimator and thus we use this estimator in all our analysis. As an extra criterion to guarantee the reliability of the estimated EW measurements, we also impose a cut-off in the difference between the Gaussian-derived and the integral estimators of less than 2 Å .

We derive the error in ΣW for single observations and how this varies with S/N from the comparison of ΣW for stars with double observations using equation (3). Figs 6(e) and (f) show that, as expected, the comparison of ΣW improves for increasing S/N . We find that the random error in ΣW from repeated measurements is well represented by $\sigma_{\Sigma W} \approx 6/(S/N)$, and hereafter we will adopt this formula as an estimate of the errors in the single measurements. This is a factor of ~ 2 larger than we estimated from theoretical calculations; however, a larger error is not surprising, given the effect of all the steps involved in the data reduction (e.g. sky subtraction and continuum estimation). Such an error in ΣW results in an $[\text{Fe}/\text{H}]$ error of ~ 0.1 dex at an $S/N \text{ per Å}$ of 20 (see Section 3).

Finally, the measurements for stars with multiple observations were combined, weighting them by their error, and this results in 1013 distinct targets for Sculptor and 1063 for Fornax. The final sample was carefully checked to weed out any spurious objects (e.g. broken fibres, background galaxies, foreground stars, etc). Excluding the spurious objects, and the objects that did not meet our S/N and velocity error criteria, our final sample of acceptable measurements consists of 648 stars in Sculptor and 944 in Fornax. Among these, two stars in Sculptor and one in Fornax did not meet our EW criteria.

As an indication of the good quality of our data, the median S/N and median error in velocity are 32.1 and 1.6 km s⁻¹, respectively, for the Sculptor data set, and 24.3 and 1.7 km s⁻¹ for the Fornax data set.

3 THE STANDARD CaT CALIBRATION WITH VLT/FLAMES USING GLOBULAR CLUSTERS

The next step is to transform the CaT EW into metallicity, [Fe/H]. The dependence of CaT line strength on metallicity is theoretically difficult to understand; however, it has been empirically proved by extensive calibration using RGB stars in globular clusters (Armandroff & Zinn 1988; Armandroff & Da Costa 1991; Olszewski et al. 1991). Rutledge et al. (1997a) presented the largest compilation of CaT EW measurements for individual RGB stars in globular clusters, which Rutledge et al. (1997b) calibrated with HR metallicities, proving the CaT method to be reliable and accurate in the range $-2.1 \lesssim [\text{Fe}/\text{H}] \lesssim -0.6$.

As summarized in Rutledge et al. (1997b), the line strength index ΣW , which is a linear combination of the EWs of individual CaT lines, depends on [Fe/H], the gravity $\log g$ and T_{eff} . As $\log g$ and T_{eff} decrease going up the RGB, it is possible to remove the effect of the gravity and temperature, taking into account the position of the star on the RGB with respect to the horizontal branch (HB).

Armandroff & Da Costa (1991), using globular clusters, showed that this is most efficiently achieved when defining a ‘reduced EW’:

$$W' = \Sigma W + \beta (V - V_{\text{HB}}), \quad (6)$$

where V_{HB} is the mean magnitude of the HB. The advantage of using $V - V_{\text{HB}}$ over, for instance, the absolute I magnitude or the $V - I$ colour is that the slope β is constant, since it does not vary with [Fe/H]. Using V_{HB} also removes any strong dependence on the distance and/or reddening. The ‘reduced EW’ W' is thus the CaT line strength index at the level of the HB. With the above definition, Armandroff & Da Costa (1991) empirically proved that W' is directly correlated with [Fe/H].

Using 52 globular clusters Rutledge et al. (1997b) found $\beta = 0.64 \pm 0.02 \text{ Å mag}^{-1}$ across the range $-2.1 \lesssim [\text{Fe}/\text{H}] \lesssim -0.6$. C04 re-determined the value of β for their sample, which included globular and open clusters. They found $\beta = 0.66 \pm 0.03 \text{ Å mag}^{-1}$ when using only globular clusters, similar to the value found by Rutledge et al. (1997b), and $\beta = 0.73 \pm 0.04 \text{ Å mag}^{-1}$ when including open clusters, which covered an higher metallicity range ($-0.6 \lesssim [\text{Fe}/\text{H}] \lesssim -0.15$).

The first uncertainty in using the CaT calibration is of course which combination of the three CaT lines should be used. In the literature, there are several examples: Rutledge et al. (1997a) used a weighted sum of the three lines; Tolstoy et al. (2001, hereafter T01) excluded the first CaT line from the sum; C04 used an unweighted sum of the three lines. Such a choice usually depends on the quality of the data set: in the case of limited S/N and with possible sky-line residual contamination, the weakest line of the CaT is usually the least reliable and so the determination of its EW is often doubtful.

These previous studies all calibrated the CaT W' on the CG97 scale, and Friel et al. (2002) for the open clusters, meaning that the relation between [Fe/H] and W' has been derived assuming that the calibration clusters have the metallicities derived by CG97. To define their [Fe/H] scale, CG97 re-analysed high-quality EWs from different sources, using a homogeneous compilation of stellar atmosphere parameters, g f values and so on, making the largest and most self-consistent analysis of this kind. It is important to note that even the HR values of the [Fe/H] for globular clusters can differ signifi-

cantly from each other (see Pritzl, Venn & Irwin 2005). Obviously, choosing a different [Fe/H] scale implies that the derived relations will be different.

This variety of approaches and calibrations can lead to a degree of confusion when viewing the literature. It is not possible to find out a priori which way of summing CaT EWs and which calibration must be used; we thus test each of the mentioned approaches from the literature and see which one performs better for our globular cluster data set. In addition, we derive our own calibration.

Each of the three calibrations that we examined from the literature consists of the following three relations.

(i) A linear combination of the CaT lines EW:

$$\Sigma W_{\text{R97}} = 0.5 \text{ EW}_1 + \text{EW}_2 + 0.6 \text{ EW}_3, \quad (7)$$

$$\Sigma W_{\text{T01}} = \text{EW}_2 + \text{EW}_3, \quad (8)$$

$$\Sigma W_{\text{C04}} = \text{EW}_1 + \text{EW}_2 + \text{EW}_3. \quad (9)$$

(ii) A relation for the ‘reduced EW’:

$$W'_{\text{R97}} = \Sigma W_{\text{R97}} + 0.64(\pm 0.02) (V - V_{\text{HB}}), \quad (10)$$

$$W'_{\text{T01}} = \Sigma W_{\text{T01}} + 0.64(\pm 0.02) (V - V_{\text{HB}}), \quad (11)$$

$$W'_{\text{C04}} = \Sigma W_{\text{C04}} + 0.73(\pm 0.04) (V - V_{\text{HB}}). \quad (12)$$

(iii) The calibration of the ‘reduced EW’ with [Fe/H]:

$$[\text{Fe}/\text{H}]_{\text{CG97}}^{\text{R97}} = -2.66(\pm 0.08) + 0.42(\pm 0.02) W'_{\text{R97}}, \quad (13)$$

$$[\text{Fe}/\text{H}]_{\text{CG97}}^{\text{T01}} = -2.66(\pm 0.08) + 0.42(\pm 0.02) W'_{\text{T01}}, \quad (14)$$

$$[\text{Fe}/\text{H}]_{\text{CG97}}^{\text{C04}} = -2.966(\pm 0.032) + 0.362(\pm 0.014) W'_{\text{C04}}, \quad (15)$$

where R97 stands for Rutledge et al. (1997a,b), T01 for Tolstoy et al. (2001) and C04 for Cole et al. (2004).

Fig. 7 shows ΣW versus $(V - V_{\text{HB}})$ for the four globular clusters, summing the CaT lines as in equations (7)–(9). As the minimum S/N per Å of these data is ~ 40 and the median is ~ 100 per Å the errors in the summed EW are very small, $\lesssim 0.1 \text{ Å}$. As a consistency check, we calculate the weighted average slope for each of the calibrations and find that they are consistent at the 1σ level with the previous work ($\beta_{\text{R97, this work}} = 0.59 \pm 0.04 \text{ Å mag}^{-1}$, $\beta_{\text{T01, this work}} = 0.62 \pm 0.03 \text{ Å mag}^{-1}$, $\beta_{\text{C04, this work}} = 0.79 \pm 0.04 \text{ Å mag}^{-1}$).

The metallicities derived from the best-fitting W' from an error-weighted linear fit of ΣW versus $(V - V_{\text{HB}})$ are summarized in Table 1. In general, the metallicities derived from these observations agree within 1σ with the metallicities on the CG97 scale. The metallicities from the C04 calibration seem to be systematically lower by ~ 0.1 dex. The T01 calibration appears to give the best performance. Thus, amongst the three relations in the literature, we will apply the T01 to our dSph data set. Fig. 8 shows that the relation between W' and $[\text{Fe}/\text{H}]_{\text{CG97}}$ derived from the four globular cluster data set is linear in each case.

In order to derive our own calibration, we should repeat all the steps, that is, derive a W' using the slope we find by fitting the summed EW versus $(V - V_{\text{HB}})$, and finding the best W' –[Fe/H] relation. However, as mentioned before, the slopes we find are consistent with those in the literature, and since the number of stars in most of our calibration globular clusters is not large, we prefer to use as relations those in the literature (equations 10–12). We just repeat the last step and find the best-fitting W' –[Fe/H] relations by performing an error-weighted linear-fit. The relations we find are

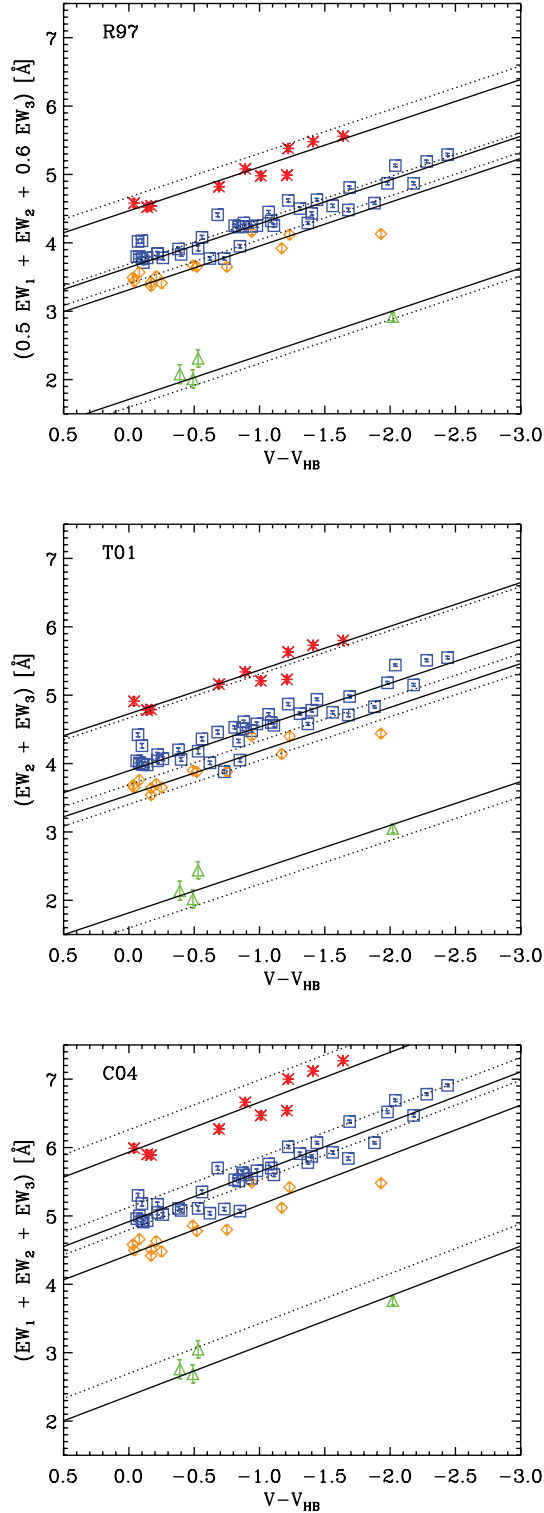


Figure 7. The CaT calibrations using globular clusters: EWs versus $(V - V_{HB})$ for RGB stars in four globular clusters of different metallicities (asterisks: NGC 104; squares: NGC 5904; diamonds: NGC 3201; triangles: NGC 4590), using three different linear combinations of CaT lines (from top to bottom panel: R97, T01, C04), using the $[\text{Fe}/\text{H}]$ published in CG97; the solid lines show our best-fitting relations, keeping the slope fixed within each calibration. The best-fitting metallicities are summarized in Table 1.

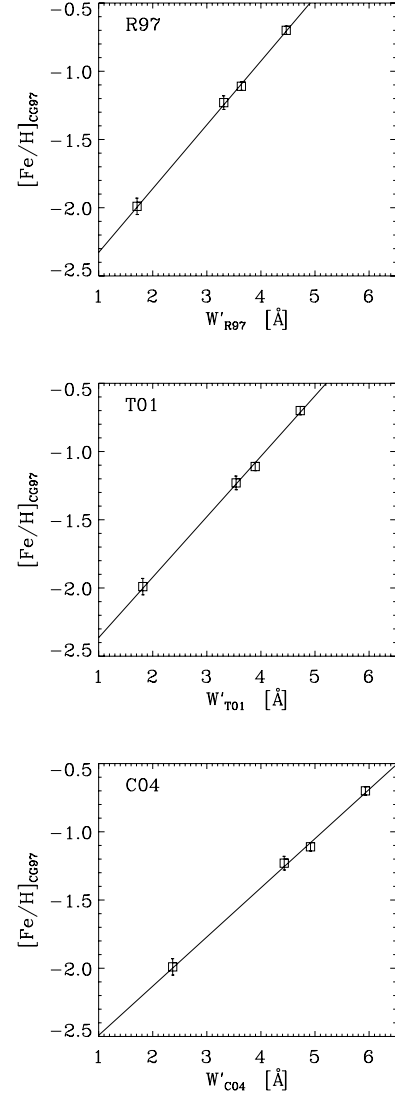


Figure 8. Relation between CG97 $[\text{Fe}/\text{H}]$ and W' for the four calibrating clusters using R97, T01 and C04 calibrations (from top to bottom panel). The solid lines show the best-fitting CG97 $[\text{Fe}/\text{H}]$ and W' relation for each calibration.

all consistent within 1σ with the relations in the literature. The calibration obtained using equation (11) appears to give the best results and is the following:

$$[\text{Fe}/\text{H}]_{\text{CG97}}^{\text{this work}} = (-2.81 \pm 0.16) + (0.44 \pm 0.04) W'_{\text{T01}}. \quad (16)$$

In the following section, we test to see if the $[\text{Fe}/\text{H}]$ – W' relation derived for the globular cluster sample is reliable when applied to RGB field stars in galaxies. To do so, we apply equations (14) and (16) to the Sculptor and Fornax samples and we compare the metallicities so derived to the HR metallicities.

4 COMPARISON TO HIGH-RESOLUTION METALLICITY MEASUREMENTS

One major uncertainty in applying the CaT method to field stars in galaxies, for example, in dSphs, is that the $[\text{Fe}/\text{H}]$ –CaT W' relations have so far been calibrated exclusively on globular clusters. Some dSphs contain intermediate-age and even young stellar populations and have a large spread in metallicity, whilst the above relations have

been derived for single-age stellar population older than 10 Gyr, over a relatively narrow metallicity range and also a very narrow range in $[\alpha/\text{Fe}]$, which is very different from composite stellar populations. Furthermore, in composite populations it is more difficult to assign a unique magnitude for the HB, although Cole, Smecker-Hane & Gallagher (2000) and C04 showed that the uncertainty due to this effect is ~ 0.05 dex, which is not significant compared to the intrinsic precision of the method (~ 0.1 dex).

The only way to reliably test the standard globular cluster calibration for dSph field stars is to compare the $[\text{Fe}/\text{H}]$ derived from CaT EWs to that obtained from direct measurements in HR observations of the same field stars. HR metallicities should be more accurate than CaT measurements because the iron abundance is not inferred from other elements but obtained by direct measurement of typically more than 60 separate Fe lines with two different ionization states (i.e. Fe I and Fe II).

Until now a comparison between HR and LR $[\text{Fe}/\text{H}]$ has not been thoroughly made for dSph field stars, partly due to the lack of a large sample of overlapping measurements. Thanks to instruments like FLAMES, it is now possible to get suitable comparison spectra for many objects at the same time.

As part of the DART Large Program at ESO, HR FLAMES spectra have been taken for 93 probable Sculptor velocity member stars in the central regions of the Sculptor dSphs (Hill et al., in preparation, expected in 2008) for which there is also LR CaT data (Tolstoy et al. 2004). A similar study was made of a central field in Fornax (Letarte et al., in preparation, expected in 2008; Letarte 2007), for which 36 stars overlap our LR sample (Battaglia et al. 2006).

These observations consist of $R \sim 20\,000$ resolution spectra of ~ 80 stars in the centre of both Fornax and Sculptor, covering ~ 60 nm in three different FLAMES set-ups (534–562 nm; 612–641 nm; 631–670 nm), and reaching typical S/N of 30 per 0.05 nm pixel. The chemical analysis of the sample was performed using OSMARCS one-dimensional stellar atmosphere models in LTE (OSMARCS models, Gustafsson et al. 2003; Gustafsson, Heiter & Edvardsson 2007), an extension of the OSMARCS models referenced above (Pletz, private communication), and a standard EW analysis. Stellar parameters were determined using a combination of photometric indices (V , I , J , H , K) and spectroscopic indicators (excitation and ionization equilibrium). The results include the abundances of ~ 10 elements, including iron and calcium which are reported here for comparison to the LR results (see Table 4). Error bars on HR abundances indicated on the plots refer to the combined abundance measurement errors and propagated stellar parameters uncertainties.

The detailed description of the data reduction and analysis of the HR spectroscopic data can be found in a series of papers (Hill et al., in preparation, expected in 2008; Letarte et al., in preparation, expected in 2008; Letarte 2007). Also note that the HR results have been put on to UVES system (e.g. Letarte 2007).

(i) *CaT calibration using HR data.* First, we determine the relation between $[\text{Fe}/\text{H}]_{\text{HR}}$ and CaT W' directly for the dSph data. We assumed $V_{\text{HB}} = 20.13$ for Sculptor and $V_{\text{HB}} = 21.29$ for Fornax, taken from Irwin & Hatzidimitriou (1995). Fig. 9 shows the HR $[\text{Fe}/\text{H}]$ of the overlapping stars between the HR and LR samples plotted against their reduced CaT EW (equation 11). The best linear fit we obtain allowing for errors in both coordinates is

$$[\text{Fe}/\text{H}]_{\text{HR}} = -2.94(\pm 0.04) + 0.49(\pm 0.01)W'. \quad (17)$$

This calibration is consistent at the 1σ level with the calibration derived in the previous section for four globular clusters; however,

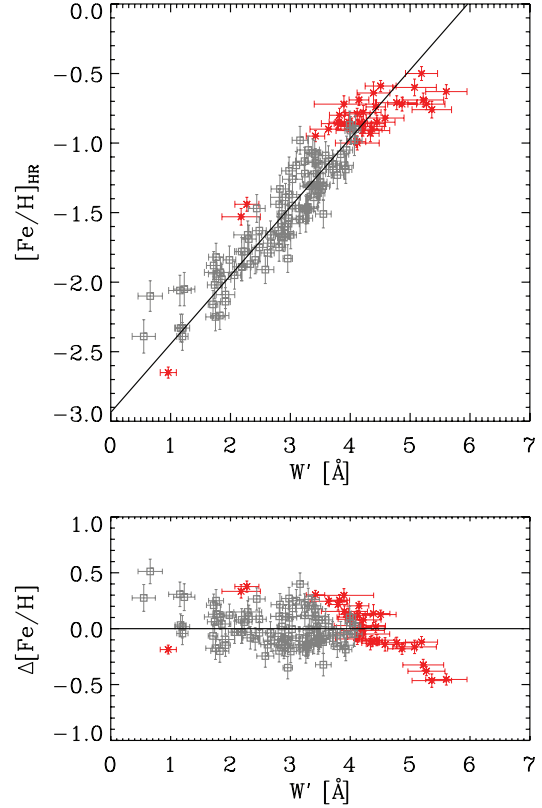


Figure 9. Top panel: HR $[\text{Fe}/\text{H}]$ versus CaT W' for the RGB stars overlapping between the HR and LR samples for the Sculptor and Fornax dSphs (93 in Sculptor, squares, and 36 stars in Fornax, asterisks). The solid line represents the best error-weighted linear fit to the data $[-2.94 (\pm 0.04) + 0.49 (\pm 0.01) W']$. Bottom panel: residuals to the linear fit. The 1σ y-error bars are the errors in HR $[\text{Fe}/\text{H}]$. The rms around the relation is 0.19 dex (MAD = 0.12 dex), comparable to the measurement errors.

there are some differences. For small W' , the above relation predicts an $[\text{Fe}/\text{H}] \sim 0.1$ dex lower than from the globular cluster calibration given in equation (16), whilst the opposite happens for large W' .

(ii) *Comparison between HR and LR results using FLAMES globular cluster calibration.* The traditional globular cluster calibration of CaT is now applied to our LR CaT W' and the results compared to $[\text{Fe}/\text{H}]_{\text{HR}}$. Fig. 10 shows the comparison between HR $[\text{Fe}/\text{H}]$ and CaT $[\text{Fe}/\text{H}]$ from our globular cluster calibration (equation 16). As also indicated in the previous figure, the two methods are generally in good agreement. The average difference is $\Delta[\text{Fe}/\text{H}] = [\text{Fe}/\text{H}]_{\text{LR}} - [\text{Fe}/\text{H}]_{\text{HR}} = -0.04 \pm 0.02$ dex and the spread is 0.17 ± 0.02 dex, which is comparable with the measurement errors. We can thus apply the globular cluster calibration to dSph field stars with some confidence, between $-2.5 \lesssim [\text{Fe}/\text{H}] \lesssim -0.8$ dex, where the relation between HR and LR data is linear. It is unclear if the comparison at $[\text{Fe}/\text{H}] > -0.8$ dex indicates a non-linearity in the relationship, as the appearance of ‘non-linearity’ is given by approximately four stars. Making a concrete statement requires more data at high $[\text{Fe}/\text{H}]$.

The average difference calculated for the entire sample would suggest the absence of systematics in our evaluation of $[\text{Fe}/\text{H}]_{\text{LR}}$. However, when plotting $\Delta[\text{Fe}/\text{H}]$ versus $[\text{Fe}/\text{H}]_{\text{HR}}$ (bottom panel, Fig. 10) a trend is visible, such that the LR $[\text{Fe}/\text{H}]$ is overestimated by ~ 0.1 dex at $[\text{Fe}/\text{H}]_{\text{HR}} \lesssim -2.2$, and underestimated to be ~ 0.1 – 0.2 dex at the high- $[\text{Fe}/\text{H}]$ end, at $[\text{Fe}/\text{H}]_{\text{HR}} \gtrsim -1.2$, which instead suggests the presence of systematics in the $[\text{Fe}/\text{H}]_{\text{LR}}$ derivation.

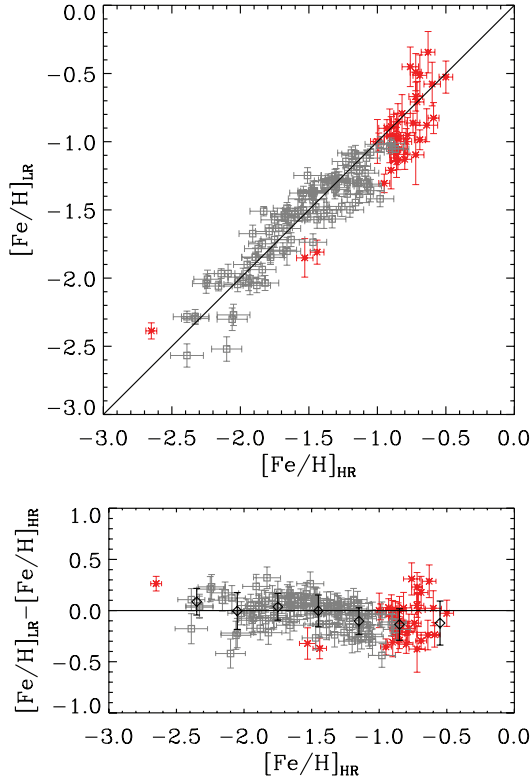


Figure 10. Top panel: comparison between HR and CaT $[\text{Fe}/\text{H}]$ for the Sculptor (squares) and Fornax (asterisks) dSphs. The solid line indicates the one-to-one relation. The CaT metallicities are derived using equation (16). Bottom panel: the above comparison plotted as a difference. The diamonds with error bars show the weighted average (and dispersion) of $[\text{Fe}/\text{H}]_{\text{LR}} - [\text{Fe}/\text{H}]_{\text{HR}}$ in bins of 0.3 dex. The average difference is $[\text{Fe}/\text{H}]_{\text{LR}} - [\text{Fe}/\text{H}]_{\text{HR}} = -0.04 \pm 0.02$ dex. The rms scatter around the one-to-one relation is 0.17 dex (MAD = 0.12 dex). For the stars with $[\text{Fe}/\text{H}]_{\text{HR}} > -0.8$, the scatter and MAD are 0.2 and 0.21 dex, respectively.

(iii) *Comparison between HR and LR results using globular cluster calibrations from the literature.* As a comparison, Fig. 11 shows the CaT $[\text{Fe}/\text{H}]$ derived from the T01 calibration (equation 14), which we used in Tolstoy et al. (2004), Battaglia et al. (2006) and Helmi et al. (2006). The trend with $[\text{Fe}/\text{H}]$ is similar to the previous case, because of the similar slopes in the $[\text{Fe}/\text{H}]$ versus W' relations. The systematic shifts are more enhanced with respect to the globular cluster calibration at the low- $[\text{Fe}/\text{H}]$ end, whilst the performance at the high- $[\text{Fe}/\text{H}]$ end is better than that for equation (16).

Both the relations in the literature and our FLAMES globular clusters calibration are in agreement with the HR data in the range $-2.5 \lesssim [\text{Fe}/\text{H}] \lesssim -0.8$ dex. Given the slight differences between the relations, the choice of one over the other does not affect considerably the derived metallicities. In the remainder of this paper, we follow the standard approach and use our own globular cluster calibration (equation 16). We also checked that $[\text{Fe}/\text{H}]$ derived from our CaT calibration is in agreement with other derivations, such as in Pont et al. (2004).

5 POSSIBLE SOURCES OF UNCERTAINTY IN THE CaT METALLICITY DETERMINATION

Although $[\text{Fe}/\text{H}]$ derived from the CaT method is in good agreement with the derivation from HR measurements, it would be interesting to understand the causes of the apparent systematics.

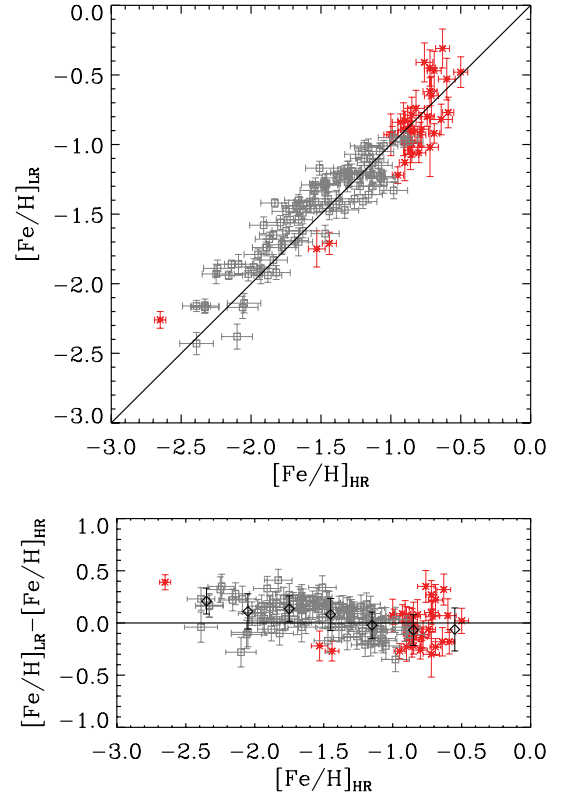


Figure 11. Top panel: the same as the previous figure, but the CaT metallicities are derived from equation (14). Though the average difference is very small, $[\text{Fe}/\text{H}]_{\text{LR}} - [\text{Fe}/\text{H}]_{\text{HR}} = 0.04 \pm 0.02$ dex, the bottom panel shows that there is a gradient in $[\text{Fe}/\text{H}]_{\text{LR}} - [\text{Fe}/\text{H}]_{\text{HR}}$ with $[\text{Fe}/\text{H}]$ more enhanced than in the previous case. The rms scatter around the one-to-one relation is 0.17 dex (MAD = 0.12 dex).

As mentioned in Section 4, the traditional CaT calibration has a number of implicit assumptions in it: (i) $[\text{Ca}/\text{H}]$ does not play a role in determining W' , and thus one can exclude it from the calibration; (ii) alternatively, if it plays a role, one assumes that globular clusters and dSphs have similar $[\text{Ca}/\text{Fe}]$; (iii) the CaT EW is a better estimator of $[\text{Fe}/\text{H}]$ than $[\text{Ca}/\text{H}]$; (iv) the effect of age differences between dSphs and globular cluster stars can be neglected. In the following, we explore the validity of these assumptions.

5.1 CaT EW as $[\text{Ca}/\text{H}]$ estimator

5.1.1 Calibration from globular clusters

To test assumption (iii), we derive a relation between W' and $[\text{Ca}/\text{H}]$ using the globular clusters as calibrators, and we apply it to our sample of dSphs. If W' traces $[\text{Ca}/\text{H}]$, then we should expect a good one-to-one relation between the Ca abundance derived from the globular cluster calibration and the Ca abundance derived from the direct HR measurements.

In order to derive the $[\text{Ca}/\text{H}]$ values needed for the globular cluster $W' - [\text{Ca}/\text{H}]$ relation, we assumed an average value of $[\text{Ca}/\text{Fe}] = 0.235$ dex for the four globular clusters (see Table 1 for the individual values), and we used $[\text{Ca}/\text{H}] = [\text{Ca}/\text{Fe}] + [\text{Fe}/\text{H}]$, where $[\text{Fe}/\text{H}]$ are the individual values from CG97 (see Table 1).

The best-fitting relation we obtain from the globular clusters is

$$[\text{Ca}/\text{H}]_{\text{CaT}} = -2.57(\pm 0.18) + 0.44(\pm 0.05) \text{ dex } \text{\AA}^{-1} W'. \quad (18)$$

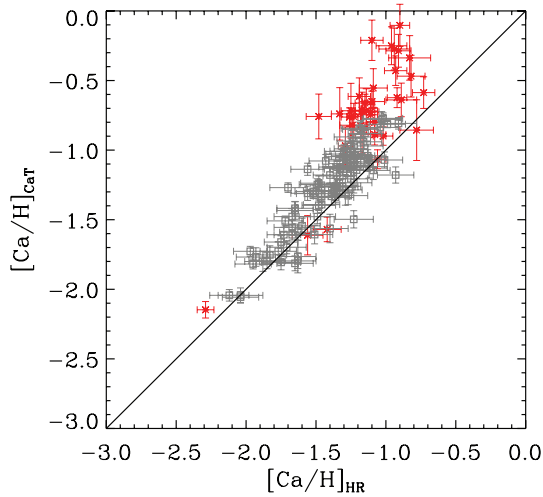


Figure 12. $[\text{Ca}/\text{H}]_{\text{CaT}}$ from equation (18) versus $[\text{Ca}/\text{H}]_{\text{HR}}$ for the Sculptor (squares) and Fornax (asterisks) dSphs. The solid line shows the one-to-one relation.

Using the individual $[\text{Ca}/\text{Fe}]$ values (instead of the average $[\text{Ca}/\text{Fe}] = 0.235$ dex) does not change the best-fitting relation, it just increases the resulting minimum χ^2 value of the fit. Fig. 12 shows the comparison between $[\text{Ca}/\text{H}]_{\text{CaT}}$ and the corresponding $[\text{Ca}/\text{H}]$ derived from HR measurements for the Sculptor and Fornax dSphs.³ The agreement between the two is not as good as for $[\text{Fe}/\text{H}]$: the scatter is much larger and $[\text{Ca}/\text{H}]_{\text{CaT}}$ is in general overestimated, especially at $[\text{Ca}/\text{H}] \gtrsim -1.3$. The comparison gets worse if we approximate the trend of $[\text{Ca}/\text{Fe}]$ with $[\text{Fe}/\text{H}]$ with a function, to circumvent the effect of the large errors in $[\text{Ca}/\text{Fe}]$.

Thus the Ca abundance cannot be the only factor that drives the W' .

5.1.2 Calibration from the literature

We compared our results to those of Bosler, Smecker-Hane & Stetson (2007). They used 15 Galactic globular clusters and six open clusters, covering the range $2 \lesssim \text{age} \lesssim 14$ Gyr and $-2.2 < [\text{Ca}/\text{H}] < +0.2$, to derive a calibration between CaT W' and Ca abundance, the idea being that using W' as a tracer of $[\text{Ca}/\text{H}]$ instead of $[\text{Fe}/\text{H}]$ one can apply the calibration to systems such as globular clusters or dSphs, without worrying about the different $[\text{Ca}/\text{Fe}]$ trends. The relation they derive is

$$[\text{Ca}/\text{H}]_{\text{Bosler}} = -2.778(\pm 0.061) + 0.470(\pm 0.016) \text{ dex } \text{\AA}^{-1} W' \quad (19)$$

and they apply it to the Leo I and Leo II dSphs. This relation is consistent with our equation (18), and this results in a poor comparison between $[\text{Ca}/\text{H}]_{\text{CaT}}$ and $[\text{Ca}/\text{H}]_{\text{HR}}$.

5.2 Exploring the effect of $[\text{Ca}/\text{Fe}]$

Another of the uncertainties in applying the CaT method to composite stellar populations (i.e. galaxies) is the varied and extended star formation histories of these systems, which results in a range of $[\text{Ca}/\text{Fe}]$ values and a trend with $[\text{Fe}/\text{H}]$. Fig. 13 shows that our

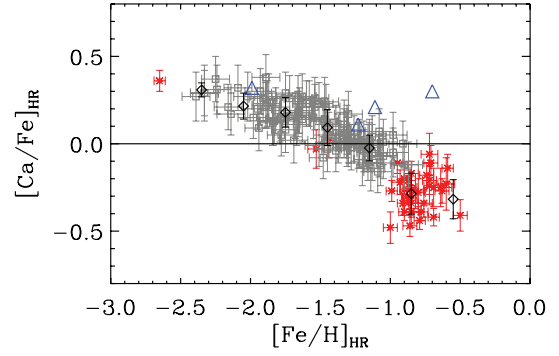


Figure 13. HR $[\text{Ca}/\text{Fe}]$ versus HR $[\text{Fe}/\text{H}]$ for the Sculptor (squares) and Fornax (asterisks) dSphs. The triangles show the values for the calibration globular clusters (Table 1). The diamonds with error bars show a weighted average and dispersion of $[\text{Ca}/\text{Fe}]$ in $[\text{Fe}/\text{H}]$ bins.

calibrating globular clusters have an almost constant $[\text{Ca}/\text{Fe}]$ trend with $[\text{Fe}/\text{H}]$ (except for NGC 3201), whilst both for the Sculptor and Fornax dSphs $[\text{Ca}/\text{Fe}]$ decreases with $[\text{Fe}/\text{H}]$. At $[\text{Fe}/\text{H}] \lesssim -1.1$ there is overlap between the $[\text{Ca}/\text{Fe}]$ values for the globular clusters and the dSphs, although $[\text{Ca}/\text{Fe}]$ in the globular clusters is larger than the average for the dSphs at the same $[\text{Fe}/\text{H}]$; at $[\text{Fe}/\text{H}] \sim -1.1$ the $[\text{Ca}/\text{Fe}]$ value for the globular clusters is ~ 0.2 dex larger than the average value for Sculptor and at $[\text{Fe}/\text{H}] = -0.7$ is ~ 0.5 larger than the average value for the Fornax dSph. This results in a smaller $[\text{Ca}/\text{H}]$ abundance for globular clusters with respect to our dSph samples at the low- $[\text{Fe}/\text{H}]$ end, and we have the opposite effect at the high- $[\text{Fe}/\text{H}]$ end. If the W' is an increasing function of both $[\text{Fe}/\text{H}]$ and $[\text{Ca}/\text{H}]$ abundances, to neglect the effect of $[\text{Ca}/\text{H}]$ in our globular cluster calibration would result in an overestimated $[\text{Fe}/\text{H}]_{\text{LR}}$ with respect to $[\text{Fe}/\text{H}]_{\text{HR}}$ in the region where $[\text{Ca}/\text{H}]_{\text{dSph}} > [\text{Ca}/\text{H}]_{\text{GC}}$; instead, we would underestimate $[\text{Fe}/\text{H}]_{\text{LR}}$ with respect to $[\text{Fe}/\text{H}]_{\text{HR}}$ in the region where $[\text{Ca}/\text{H}]_{\text{dSph}} < [\text{Ca}/\text{H}]_{\text{GC}}$. This goes in the same direction as what we see in Fig. 10. This suggests that $[\text{Ca}/\text{H}]$ might also play a role in determining the CaT W' .

It is not obvious how to quantify this effect as it is not known how the CaT W' depends on $[\text{Ca}/\text{H}]$ (linearly, quadratically, etc). We tested for a simple dependence of this kind:

$$W' = a[\text{Fe}/\text{H}] + b[\text{Ca}/\text{H}] + c, \quad (20)$$

where a , b and c are constants. We derive $a = 1.58 \pm 0.09$, $b = 0.36 \pm 0.13$ and $c = 5.85 \pm 0.07$ by fitting the observed values of W' from the dSph data set as a function of the corresponding $[\text{Fe}/\text{H}]$ and $[\text{Ca}/\text{H}]$ from HR measurements. This relation, which suggests that $[\text{Ca}/\text{H}]$ does have an impact on W' , but it is less significant than $[\text{Fe}/\text{H}]$, agrees well with the observed values of W' when applied to the globular cluster data (Fig. 14), although it does not remove the trend altogether. The importance of $[\text{Ca}/\text{H}]$ in driving W' is anyway unclear since the comparison between the observed W' for the globular clusters and the W' predicted by applying equation (17) – therefore neglecting the effect of $[\text{Ca}/\text{H}]$ – improves slightly at the low- W' end while gets slightly worse at intermediate W' .

This analysis is clearly not exhaustive, but indicative of the effects of the obvious sources of uncertainty (see also Cenarro et al. 2002). Since we are interested in understanding the effect on the $[\text{Fe}/\text{H}]$ determination, it is important to note that notwithstanding the large difference in $[\text{Ca}/\text{Fe}]$ between calibrating globular clusters and dSphs at the high- $[\text{Fe}/\text{H}]$ end (~ 0.5 dex at $[\text{Fe}/\text{H}] \sim -0.7$) and the large difference in age (Fornax stars at ~ -0.7 have an age of

³ Out of the 93 stars in the overlapping sample in Sculptor, four had too low S/N to allow a determination of the HR $[\text{Ca}/\text{H}]$.

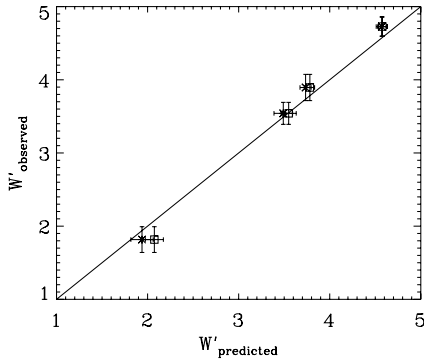


Figure 14. Observed CaT W' for our globular cluster sample versus CaT W' derived from a linear combination of Fe and Ca abundances (equation 20), calibrated on the dSph sample (squares) and from the linear dependence on Fe abundance given by equation (17) (asterisks). The solid line indicates the one-to-one relation.

3–6 Gyr) the error in estimating $[\text{Fe}/\text{H}]$ using globular clusters as calibrators is just 0.1–0.2 dex.

We point out that in this respect the Fornax dSph is one of the most extreme cases as the large majority of dSphs do not show such an extended star formation history, and so the difference in $[\text{Ca}/\text{Fe}]$ (and in age) with respect to the calibrating globular clusters will typically not be as extreme.

5.3 On age effects

As shown in Pont et al. (2004), for metallicities < -0.5 dex, a relative age difference of ~ 12 Gyr with respect to calibrating globular clusters results in an underestimate in metallicity of ~ 0.4 dex (their fig. 16, middle panel). This error decreases substantially for smaller age differences. For instance, for a metallicity ~ -0.5 , the error induced by using a calibrating cluster 12.6 Gyr old for a star which is 4 Gyr old is $\sim \pm 0.2$ dex. At lower metallicities, this effect is smaller. The stars in the Sculptor dSph have ages > 8 Gyr (e.g. Monikiewicz et al. 1999) and $[\text{Fe}/\text{H}] \lesssim -1$, thus the age effect will be negligible. The bulk of stars in the Fornax dSph have an age between 3 and 6 Gyr old (e.g. Pont et al. 2004; Battaglia et al. 2006); only a small fraction of stars have younger ages,⁴ thus the age effect could result in an underestimate of ~ 0.1 – 0.15 dex from the CaT measurement for 3–6 Gyr old stars. As mentioned above, most of the dSphs do not show star formation histories as extended as Fornax, thus the age effect should typically be even smaller.

5.4 Synthetic stellar atmospheres

Another way of approaching these issues is to model the expected behaviour of the CaT lines using stellar atmospheres and synthetic spectra. It is perhaps not surprising that these lines exhibit a strong dependence on global metallicity, best represented by $[\text{Fe}/\text{H}]$, rather than on the abundance of calcium itself, given the strength of the features. With EWs ranging from ~ 500 mÅ to several Å, the CaT lines are not only saturated but also truly in the strong line regime of line formation, when a significant part of the flux is found in the pressure-broadened wings rather than the core of the line. This effectively means that electronic pressure (P_e) is driving the strength

of the line rather than the usual combination of temperature and elemental abundance that shape weak lines. This can explain a fundamental characteristic of the CaT: since many metals contribute to P_e , the CaT lines become sensitive to the global metallicity, $[\text{Fe}/\text{H}]$, rather than calcium abundance alone, through the pressure-dependent wings. In fact, natural broadening dominates the wings of CaT lines, so that line strength will increase with decreasing electronic pressure (roughly as $1/P_e$). This explains why the CaT line strengths increase with increasing luminosity (i.e. decreasing gravity and hence pressure). The metallicity dependence of P_e will therefore also contribute to shaping the CaT lines. Finally, global metallicity also plays a role in changing the blanketing properties by numerous small metallic absorption lines, both in the wavelength regions used to define the continuum and in the CaT wings themselves. It is therefore difficult to speculate theoretically how the CaT W' should behave upon varying each stellar parameter and we have therefore used synthetic stellar spectra, ‘observed’ with an approach mimicking our observational procedure.

We have checked the validity of calibrating $[\text{Fe}/\text{H}]$ to the width of the CaT lines by using synthetic spectra of RGB stars. Munari et al. (2005) published a large grid of synthetic spectra covering the CaT region at various resolutions. Since our FLAMES LR of $R \approx 6500$ did not correspond to the available resolutions, we rebinned the Munari $R = 20\,000$ resolution spectra to the FLAMES LR. The model atmosphere spectra cover a range of stellar parameters including $\log g$, T_{eff} and $[\text{M}/\text{H}]$ and also include models computed with $[\alpha/\text{Fe}] = 0.0$ and $+0.4$. The CaT EWs of a representative sample of model atmosphere spectra covering the range of stellar RGB parameters encountered in dSphs were measured in a similar manner to those of our observed LR spectra.

In particular, we measured the EW of a set of synthetic stars taken along the upper RGB of a series of 12-Gyr isochrones with $-2.5 \leq [\text{Fe}/\text{H}] \leq +0.0$ (Bertelli et al. 1994). We then derived the corresponding $[\text{Fe}/\text{H}]$ using our CaT calibration based on globular clusters (equation 16). The V magnitude for each synthetic star was read off the isochrone, while the V_{HB} was assumed $[\text{Fe}/\text{H}]$ -dependent, of the form $V_{\text{HB}} = 1.17 + 0.39[\text{Fe}/\text{H}]$ (Ortolani et al. 1995, and references therein, which also fit well the same isochrones). The resulting $(V - V_{\text{HB}})$ range from 0.5 to 3.5 (RGB tip), and are very similar to the range of luminosities of our targets in Sculptor.

The CaT EWs were measured in both α -poor and α -rich models, and Fig. 15 shows how remarkably well the resulting metallicities from the CaT calibration ($[\text{Fe}/\text{H}]_{\text{CaT}}$) agree with the input model metallicity ($[\text{Fe}/\text{H}]_{\text{model}}$). Our calibration on the α -rich synthetic spectra overestimates the $[\text{Fe}/\text{H}]$ model by only $+0.19$ dex with a negligible dispersion and no particular trend, while our $[\text{Fe}/\text{H}]_{\text{CaT}}$ of the α -poor spectra underestimates the $[\text{Fe}/\text{H}]$ model by -0.38 dex in the mean. This is expected, as our calibration is based on globular clusters, which have in the mean $[\alpha/\text{Fe}] \sim +0.25$, and are therefore closer to the α -rich grid of Munari et al. (2005). Moreover, the relation between the CaT W' and $[\text{Fe}/\text{H}]$ model is linear, with a slope (for α -rich) matching exactly the slope of our calibration. This reassuring result supports the soundness of using simple CaT (linear) calibrations to derive metallicities.

We also examined the effect of using the same calibration for synthetic stars along younger RGB isochrones of 2.5 Gyr (similar to the mean age of the stars in our Fornax sample). The main effect is the luminosity increase: at a given metallicity, the RGB isochrones of intermediate and old ages overlap almost perfectly in temperature and gravity, but younger RGB stars appear more luminous than old stars of the same temperature and gravity (because a younger star with the same gravity and temperature will be of higher mass, and

⁴ These spectroscopic samples of RGB stars only contain stars older than 1–2 Gyr.

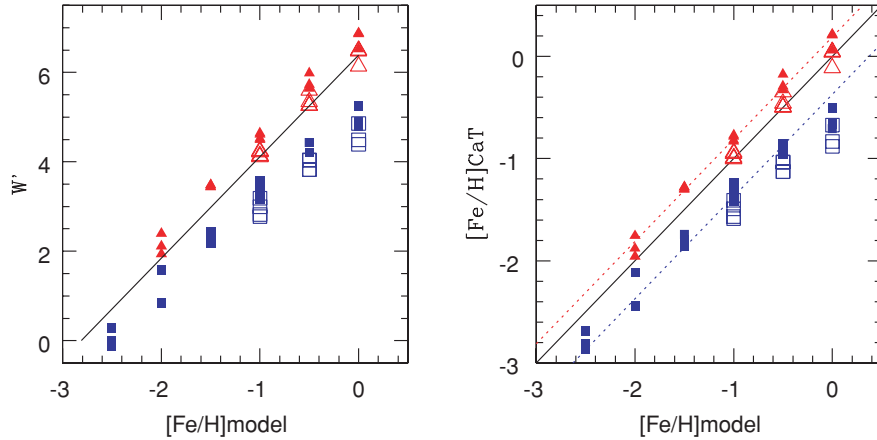


Figure 15. Left-hand panel: relation between the W' and the input metallicity of the synthetic spectra of the Munari et al. (2005) grid ($[Fe/H]_{\text{model}}$) along a 12-Gyr isochrone. The filled squares are the models with $[\alpha/Fe] = 0.0$ and the filled triangles are the models with $[\alpha/Fe] = 0.4$. The thick black line shows our calibration. The open symbols show the effect of changing the age of the isochrone from 12 to 2.5 Gyr for metallicities $[Fe/H] \geq -1$. Right-hand panel: comparison between the metallicity obtained using our CaT calibration on the synthetic spectra ($[Fe/H]_{\text{CaT}}$) and the input metallicity of the model ($[Fe/H]_{\text{model}}$). The dotted lines show the mean offsets of the $[\alpha/Fe] = 0.0$ models and $[\alpha/Fe] = 0.4$ models with respect to the one-to-one relation (solid line).

hence more luminous). Between a 12- and 2.5-Gyr isochrone, the luminosity at a given gravity increases by ~ 0.5 mag, translating into a -0.2 dex metallicity decrease once the CaT calibration is applied, as illustrated in Fig. 15. This age effect is small, and comparable to the shallow slope observed in Fig. 10 (lower panel) where our younger metal-rich stars (Fornax) tend to give lower $[Fe/H]$ CaT abundances than expected.

As a final comment on this comparison, we would like to stress that the CaT behaviour as a function of luminosity and metallicity assumed in this work is fully consistent with what is expected from synthetic spectra. The synthetic spectra are, however, probably not precise enough to further constrain the CaT behaviour, as the core of these very strong lines is almost certainly not well modelled by the stellar atmosphere models, as the cores of the lines form close to the stellar surface, a region which is problematic to model, and may even include a possible contribution from the stellar chromosphere.

6 SUMMARY AND CONCLUSIONS

We described the data-reduction steps we use within the DART collaboration to estimate velocities and CaT metallicities for LR data for RGB stars in dSphs. We showed that we obtain accurate velocities and $[Fe/H]$ measurements, with internal error in velocity $\sim 2 \text{ km s}^{-1}$ and in $[Fe/H] \sim 0.1$ dex at an S/N per \AA of 20.

We used four Galactic globular clusters observed with VLT/FLAMES in the CaT region to test the performance of several CaT W' – $[Fe/H]$ relations existing in the literature. We also derived the best calibration from these globular cluster data. The relation here derived is consistent at the 1σ level with the calibration derived in T01, which we used in Tolstoy et al. (2004), Battaglia et al. (2006) and Helmi et al. (2006).

We used a sample of 93 and 36 RGB stars in the Sculptor and Fornax dSphs, respectively, overlapping between our LR and HR VLT/FLAMES observations, to test the globular cluster CaT calibration. This is for the first time that the CaT calibration is tested on field stars in galaxies. We find a good agreement between the metallicities derived with these two methods. However, a systematic trend is present with $[Fe/H]$, such that using the globular cluster calibration derived in this work the $[Fe/H]$ measurement from CaT is overestimated by ~ 0.1 dex at $[Fe/H]_{\text{HR}} < -2.2$, whilst at

$[Fe/H]_{\text{HR}} > -1.2$ it is underestimated by ~ 0.1 – 0.2 dex. No clear systematic trend is instead derived from our data for $[Fe/H]_{\text{HR}} > -0.8$. In order to understand this systematic effect, we explored the possible contribution of Ca abundance to the calibration, and showed that there are indications that it might well affect the CaT W' , although much less than $[Fe/H]$. From our data set, we also show that, contrary to previous claims, it is not advisable to use the CaT W' as a linear indicator of $[Ca/H]$.

Finally, we investigated the effect of varying stellar atmosphere parameters on the CaT method by analysing a large sample of model atmosphere spectra (Munari et al. 2005). We again demonstrated that the CaT method is (surprisingly) robust to the usual combination of age, metallicities and $[\alpha/Fe]$ variations seen in nearby dSphs.

From our analysis, we see that even for large differences in $[Ca/Fe]$ between calibrating globular clusters and our sample of dSphs (~ 0.5 dex at $[Fe/H] \sim -0.7$) and large difference in ages (at $[Fe/H] \sim -0.7$ Fornax stars are ~ 10 Gyr younger than globular cluster stars) the error in estimating $[Fe/H]$ using globular clusters as calibrators is just 0.1 – 0.2 dex. The Fornax dSph is likely to represent the most extreme case as it has had one of the most extended star formation histories among the dSphs in the Local Group.

We conclude that CaT– $[Fe/H]$ relations calibrated on globular clusters can be applied with confidence to RGB stars in composite stellar populations such as galaxies, at least in the $[Fe/H]$ range probed by the above analyses, $-2.5 < [Fe/H] < -0.5$. Hence, the CaT method provides a good indicator of the overall metallicity of resolved stars.

This has implications for the efficiency with which we can obtain metallicity distribution functions of nearby resolved galaxies. The HR data collected in this paper required more than six nights of VLT observing time for ~ 150 spectra, whereas ~ 120 CaT spectra were obtained in 1-h VLT observing time.

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SUPPLEMENTARY MATERIAL

The following supplementary material is available for this article:

Table 3. This table lists the relevant data for the stars in the globular clusters NGC 104, NGC 3201, NGC 4590 and NGC 5904 observed with VLT/FLAMES at LR, and that we used in the analysis of the CaT–[Fe/H] calibration. For the CaT analysis, we select only those stars whose colour and magnitude are consistent with what is expected for RGB stars, and that have $V - V_{\text{HB}} < 0$, $S/N > 10$ per Å error in velocity $< 5 \text{ km s}^{-1}$ and velocities consistent with membership to the cluster (we assign membership to those stars within 3σ of the systemic velocity, see Table 1). The columns indicates: (1): the globular cluster name; (2): the star ID; (3) and (4): star co-ordinates (right ascension in hours and declination in degrees); (5) and (6): V magnitude and its error; (7) and (8): heliocentric velocity and its error; (9) and (10): summed CaT EW ($\text{EW}_2 + \text{EW}_3$) and its error; (11) and (12): [Fe/H] value and its error, derived from the LR observations applying equation (16). The photometry and astrometry are from Stetson (2000).

Table 4. This table lists the relevant data for the stars observed with VLT/FLAMES at both LR and HR in the Sculptor and Fornax dSphs, and used in the analysis of the CaT–[Fe/H] calibration. The columns indicates: (1): the galaxy name; (2): the star ID; (3) and (4): star coordinates (right ascension in hours and declination in degrees); (5) and (6): V magnitude and its error; (7) and (8): heliocentric velocity and its error; (9) and (10): summed CaT EW ($\text{EW}_2 + \text{EW}_3$) and its error; (11) and (12): [Fe/H] value derived from the LR observations applying equation (16) and its error; (13) and (14): HR [Fe/H] value and error; (15) and (16): [Ca/Fe] and its error, from the HR observations. The photometry and astrometry are from our ESO/WFI observations (Tolstoy et al. 2004; Battaglia et al. 2006). Those stars with [Ca/Fe] = -9.99 had too low S/N to allow for a determination of the HR [Ca/H].

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